

REST-FRAME OPTICAL SPECTRA AND BLACK HOLE MASSES OF $3 < Z < 6$ QUASARS

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ABSTRACT

We present the rest-frame optical spectral properties of 155 luminous quasars at $3.3 < z < 6.4$ taken with the *AKARI* space telescope, including the first detection of H α emission line as far out as $z \sim 6$. We extend the scaling relation between the rest-frame optical continuum and line luminosity of active galactic nuclei (AGNs) to the high luminosity, high redshift regime that has rarely been probed before. Remarkably, we find that a single log-linear relation can be applied to the 5100 Å and H α AGN luminosities over a wide range of luminosity ($10^{42} < L_{5100} < 10^{47}$ ergs s⁻¹) or redshift ($0 < z < 6$), suggesting that the physical mechanism governing this relation is unchanged from $z = 0$ to 6, over five decades in luminosity. Similar scaling relations are found between the optical and the UV continuum luminosities or line widths. Applying the scaling relations to the H β black hole mass (M_{BH}) estimator of local AGNs, we derive the M_{BH} estimators based on H α , Mg II, and C IV lines, finding that the UV-line based masses are overall consistent with the Balmer-line based, but with a large intrinsic scatter of 0.40 dex for the C IV estimates. Our 43 M_{BH} estimates from H α confirm the existence of BHs as massive as $\sim 10^{10} M_{\odot}$ out to $z \sim 5$, and provide a secure footing for previous Mg II-line based studies that a rapid M_{BH} growth has occurred in the early universe.

Subject headings: galaxies: active — galaxies: evolution — quasars: emission lines — quasars: super-massive black holes

1. INTRODUCTION

Quasars, galaxies that are in an active phase due to vigorous accretion of matter toward the central supermassive black hole (BH), has been vastly discovered by many surveys (e.g., Schmidt & Green 1983; Hewett et al. 1995; Boyle et al. 2000; York et al. 2000; Richards et al. 2002; Im et al. 2007; Lee et al. 2008; Willott et al. 2010b; Wu et al. 2010). Through the discovery of

quasars at high redshift (Fan et al. 2000; Cool et al. 2006; Goto 2006; Stern et al. 2007; Willott et al. 2007; Mortlock et al. 2011; Venemans et al. 2013; Bañados et al. 2014), we are witnessing the early stages of supermassive BH growth in the distant universe. The number density of optically luminous quasars at high redshift, quickly increases with cosmic time towards its maximum at $z = 2 \sim 3$ (Dunlop & Peacock 1990; Warren et al. 1994; Schmidt et al. 1995; Kenefick et al. 1995; Richards et al. 2006b; McGreer et al. 2013). Accompanied by high accretion rates among luminous quasars at $z > 4$ (e.g., Willott et al. 2010a; De Rosa et al. 2011; De Rosa et al. 2014), this suggests a rapid BH growth of the luminous population of active galactic nuclei (AGNs) in the early universe. Also, an unusual population of AGNs known as dust-poor quasars – quasars with little infrared emission from hot and warm dust – is found to be more common at higher redshift (Jiang et al. 2010; Jun & Im 2013; Leipski et al. 2014). Luminous dust-poor quasars tend to have lower BH masses (M_{BH}) or higher Eddington ratios compared to typical luminous quasars (Jiang et al. 2010; Jun & Im 2013), indicating the build-up of the AGN dusty substructure during its early mass accretion.

One of the key findings in the study of high redshift quasars is that there exist extremely massive BHs, with the mass reaching $M_{\text{BH}} \sim 10^{10} M_{\odot}$ at $z = 2\text{--}5$, and $\sim 10^9 M_{\odot}$ at $z = 6\text{--}7$ (Jiang et al. 2007; Netzer et al. 2007; Kurk et al. 2007; Vestergaard et al. 2008; Shen et al. 2008; Mortlock et al. 2011; De Rosa et al. 2014). Under the concordance cosmology, the time gap between the reionization epoch of the universe from the recent *Planck* study of the cosmic microwave background, $z \sim 11.5$ (Planck Collaboration et al. 2014), and $z = 6$ is 0.5 Gyr.

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Considering the case where a Population II stellar seed M_{BH} starts to grow at $z \sim 11.5$, the given time is too short for the seed to become an extremely massive BH at $z=6$. Under the Eddington-limited accretion where the mass accretes at a maximal rate with the radiative pressure gradient and gravity in balance, we expect a BH to grow as

$$M(t) = M_0 \exp\left(\frac{1 - \epsilon}{\epsilon} \frac{t - t_0}{t_{\text{Edd}}}\right), \quad (1)$$

where M_0 is the mass at t_0 , ϵ is the radiative efficiency, and t_{Edd} is the Eddington limited timescale of 0.45 Gyr. Assuming a typical value of $\epsilon = 0.1$, a BH can grow by 2×10^4 times over the time span of 0.5 Gyr, without considering feedback mechanisms that could slow down the BH growth. This maximal growth factor is far too small for a stellar mass BH with a typical seed mass of $10 M_{\odot}$, to grow into the extremely massive AGNs that have been observed recently. Consequently, BH seeds that may have started accreting prior to the reionization epoch, which are more massive (e.g., Bromm & Loeb 2003; Begelman et al. 2006; Lodato & Natarajan 2006; Bellovary et al. 2011) or go through super-Eddington accretion (e.g., Volonteri & Rees 2005; Wyithe & Loeb 2012; Madau et al. 2014), are suggested to explain the M_{BH} of quasars at high redshift (also see a review on this subject by Natarajan 2014).

Obviously, accurate determination of M_{BH} is an important requirement for understanding the BH growth at high redshift. This is especially true for BHs at the most massive end. A large uncertainty in M_{BH} can scatter the abundant lower mass BHs into the high mass end of M_{BH} distribution, while the effect in the opposite direction is much less significant since higher mass BHs are relatively rare. As a result, the number of extremely massive BHs can be easily overestimated. In principle, the correction to this effect is possible, but it requires a good knowledge on the error of M_{BH} measurements which is, however, rather difficult to obtain. This poses a potential challenge to the understanding of the BH growth at high redshift as described below.

In most of high redshift quasar studies, BH masses are estimated using the rest-frame ultraviolet (UV)-part of spectra which is redshifted into the rest-frame optical (rest-optical). The velocity widths of broad UV lines such as C IV and Mg II are used as measures of the gas motion of the broad line region (BLR), and the continuum or the line luminosity at the rest-frame UV (rest-UV) is used as a proxy for the size of BLR (R_{BLR}). One gets M_{BH} by combining the two pieces of information through a virial mass estimator, $M_{\text{BH}} \propto R_{\text{BLR}} \times \text{FWHM}_{\text{BLR}}^2$ (e.g., McLure & Dunlop 2004; Vestergaard 2004; Baskin & Laor 2005; Sulentic et al. 2007; Shen et al. 2008; Park et al. 2013). While UV-based M_{BH} estimators are useful tools to measure the M_{BH} of AGNs, they are secondary estimators that are derived from rest-optical spectral properties such as the line luminosity or width of $\text{H}\beta$ and $\text{H}\alpha$, and the continuum luminosity at 5100 \AA (L_{5100}). Assuming the UV luminosity follows the optical broad line region radius–luminosity ($R_{\text{BLR}}-L$) relation with a constant factor, and the UV broad line width follows the optical line width as a power law relation, the UV M_{BH} estimators are derived. Con-

sequently, a number of studies have been carried out to justify the use of Mg II or C IV-line based M_{BH} estimators comparing the masses from UV estimators to those from the optical. While some studies suggest that UV-line M_{BH} estimators are reasonably accurate, especially for Mg II, other studies point out a large scatter between C IV-based measurements versus $\text{H}\beta$ -based measurements which can make the C IV-based M_{BH} values uncertain by a factor of a few (e.g., Netzer et al. 2007; Shen & Liu 2012, hereafter S12). It has been noted that non-virialized motion (e.g., Denney 2012) or extinction (e.g., Assef et al. 2011), could severely modify the C IV line profile such that the M_{BH} cannot be reliably measured from a simple virial equation, although it is controversial on the exact origin of the discrepant UV-based M_{BH} with that of optical.

Furthermore, for the M_{BH} estimates to be valid, one also needs to justify the application of the low redshift M_{BH} estimators to high redshift, luminous quasars. Although one can expect that the M_{BH} estimators should not evolve in time in any significant way based on physical ingredients of AGN models, this has not been tested at the high redshift, high luminosity regime. An ultimate test would be to perform a reverberation mapping study of high luminosity quasars at high redshift, but such a study would take decades to complete, since the variability timescale is long for luminous quasars and the cosmological time dilation makes it even longer. Another way, albeit less direct than the reverberation mapping method, would be to investigate the correlation between the line and continuum luminosities. At low redshift, the $\text{H}\alpha$ or $\text{H}\beta$ line luminosities are known to tightly correlate with the optical continuum luminosity (e.g., Greene & Ho 2005). As the radiation energy L from an accretion disk increases, the distance to the broad line region increases as $R \propto L^{0.5}$ from a simple photo-ionization argument, or the energy flux of the radiation incident upon BLR clouds would be independent of the luminosity of the central power source, which has been confirmed observationally (Kaspi et al. 2000; Kaspi et al. 2005; Bentz et al. 2013). A modification in the $R_{\text{BLR}}-L$ relation at high redshift or high luminosity for example, would result in an increase or a decrease in the incident energy flux upon the BLR, thus the correlation between the line and continuum luminosities is likely to be modified. Interestingly, several studies suggest that the $R_{\text{BLR}}-L$ relation is not valid for luminous AGNs having massive black holes with low spins (Laor & Davis 2011; Wang et al. 2014), due to a decrease in the ionizing flux L_{ion} caused by a decrease in the radiation temperature of the accretion disk. Since the line luminosity L_{line} is proportional to L_{ion} , such theoretical expectations can be tested by examining the $L_{\text{line}}-L_{5100}$ relation at luminous end.

In order to estimate the M_{BH} using an optical mass estimator and to test the universality of the scaling relations in the rest-optical for quasars at $z > 3.5$, spectroscopic observation is necessary at $\lambda > 2.5 \mu\text{m}$. This however, is a very challenging task from the ground due to high thermal background at $\lambda > 2.5 \mu\text{m}$ and atmospheric absorptions, limiting such efforts to the study of $\text{H}\beta$ line at $z < 3.5$ (e.g., Shemmer et al. 2004; Netzer et al. 2007; Assef et al. 2011). Recently, *AKARI* spectroscopic observations have provided a breakthrough for the

study of the rest-optical spectra of distant objects, where its unique 2.5–5.0 μm coverage enables the redshifted $\text{H}\alpha$ line to be probed from $z=3$ to 6.5. Oyabu et al. (2007) reported the *AKARI* detection of the redshifted $\text{H}\alpha$ line of a quasar RX J1759.4+6638 at $z=4.3$, and Sedgwick et al. (2013) investigated the star formation rate of submm galaxies at $z > 3.5$ based on $\text{H}\alpha$.

With one of the *AKARI* mission programs (guaranteed time) and also through several small open time programs, we performed a spectroscopic study of quasars at high and low redshifts, with the aim to obtain the rest-optical spectra of high redshift quasars or the rest-frame 2.5–5.0 μm spectra of low redshift AGNs. We call all these programs QSONG (Quasar Spectroscopic Observations with NIR Grism) after the name of the mission program, and here, we present the rest-optical spectral properties of 155 type-1 quasars from QSONG, along with scaling relations and M_{BH} estimates based on these spectra.

The contents of the paper are constructed as follows. First, we use the spectro-photometric data including *AKARI* for $z \gtrsim 3$ AGNs (section 2), and derive the continuum and line emission properties (section 3) in order to calibrate the $\text{H}\alpha$ M_{BH} for the usage in high redshift. We check the validity of the continuum and broad line luminosity relations and line width relations with respect to the local, update the mass equations for $\text{H}\alpha$, $\text{H}\beta$, Mg II , and C IV , and compare the UV–optical M_{BH} 's (section 4). Finally, we discuss on the reliability of single epoch M_{BH} estimators at high redshift, and investigate the massive end evolution of M_{BH} in distant AGNs (section 5). Throughout this paper we adopt a flat ΛCDM cosmology with parameters of $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$. For the virial factor in the M_{BH} estimator, we adopt $f = 5.1 \pm 1.3$ based on the calibration of the $M_{\text{BH}}-\sigma$ relation using the combined sample of AGNs and quiescent galaxies by Woo et al. (2013).

2. DATA

2.1. Sample

The majority of the data comes from the *AKARI* program QSONG, which is a two-year warm (phase-3) mission program consisting of ~ 900 approved pointings, or ~ 150 hours of observations. The program is aimed to obtain the rest-optical spectra of high redshift AGNs (Jun et al. 2012) or 2.5–5.0 μm spectra of low redshift AGNs (Kim et al. 2015) containing Brackett and 3.3 μm Polycyclic aromatic hydrocarbon lines. The sample for this study is limited to the high redshift, occupying 65% of the entire QSONG data. It is composed of optically luminous and spectroscopically confirmed type-1 quasars at $z \gtrsim 3$, mostly out of SDSS (DR5 catalog, Schneider et al. 2007; and additional discoveries from Fan et al. 2000; Fan et al. 2001; Fan et al. 2003; Fan et al. 2004; Fan et al. 2006) and APM-UKST (Storrie-Lombardi et al. 1996; Storrie-Lombardi et al. 2001) surveys. Additional targets are from various references (Webb et al. 1988; Gregory & Condon 1991; Henry et al. 1994; Griffith et al. 1995; Kennefick et al. 1995; Djorgovski et al. 1998; Rengstorf et al. 2004; McGreer et al. 2006).

The targets are type-1 AGNs so that they allow the

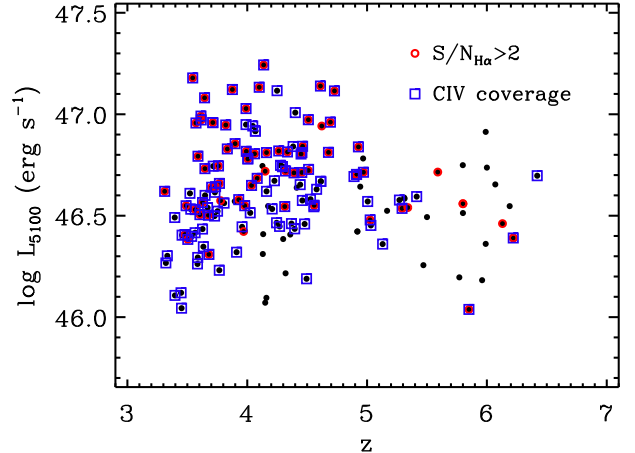


FIG. 1.— 5100 Å luminosity–redshift distribution of our *AKARI* observed quasars (black dots). The subsample with $S/N_{\text{H}\alpha} > 2$ are pointed in red, and objects with rest-UV spectral coverage including the C IV emission are marked as blue squares.

M_{BH} estimation from the broad line kinematics, through $\text{H}\alpha/\text{H}\beta$ appearing within the *AKARI* near-IR (NIR) spectral coverage. In order to provide a minimal sensitivity limit to the sample, we first considered the aperture size (68 cm) of the telescope and the restricted exposure time available at each sky position. After simulating the rest-optical spectra under the expected *AKARI* signal-to-noise ratio (S/N), the targets were chosen with z -band flux limits of ~ 18.5 and 19 AB magnitudes for bright and faint subsamples, respectively, with longer exposure time assigned to fainter targets for a clear line detection. Moreover, the targets were bounded in $3.3 < z < 6.4$ so that the $\text{H}\alpha$ emission and the surrounding continuum are placed within the 2.5–5.0 μm window of *AKARI* NIR spectroscopy. Without further constraints the targets were randomly selected in coordinates, redshift, and luminosity. We plot the distribution of $z-L_{5100\text{\AA}}$ in Figure 1.

Following the $\text{H}\alpha$ observations of distant AGNs from *AKARI* NIR spectroscopy (Oyabu et al. 2007; Oyabu et al. 2009), our initial $\text{H}\alpha$ survey of 14 quasars at $z \sim 6$ under the Helium-cooled (phase-2) open time program HZQSO (Im 2010), demonstrated the feasibility of *AKARI* observations in detecting the redshifted $\text{H}\alpha$ emission. QSONG is essentially a phase-3 extension of the survey, for the purpose of vastly expanding the number of targets at the expense of warm phase sensitivity. Thus it probes a lower redshift distribution of quasars than HZQSO, with a peak at $z \sim 4$. In addition, two more phase-3 open time programs, HQSO2 and DPQSO, were carried out either to push the redshift limit of QSONG or to detect fainter optical lines ($\text{H}\beta$ and $[\text{O III}]$), from deep exposures. Unfortunately, the Helium-dry observations led to significantly higher noise levels than expected, restricting the distinct scientific goals of the phase-3 programs that required better sensitivity. Therefore, we decided to merge all open time programs listed above under the scope of QSONG.

2.2. Data acquisition

We mostly used the NIR grism (NG) mode of the Infrared Camera (IRC, Onaka et al. 2007; Ohya et al. 2007) onboard the *AKARI* satellite (Murakami et al.

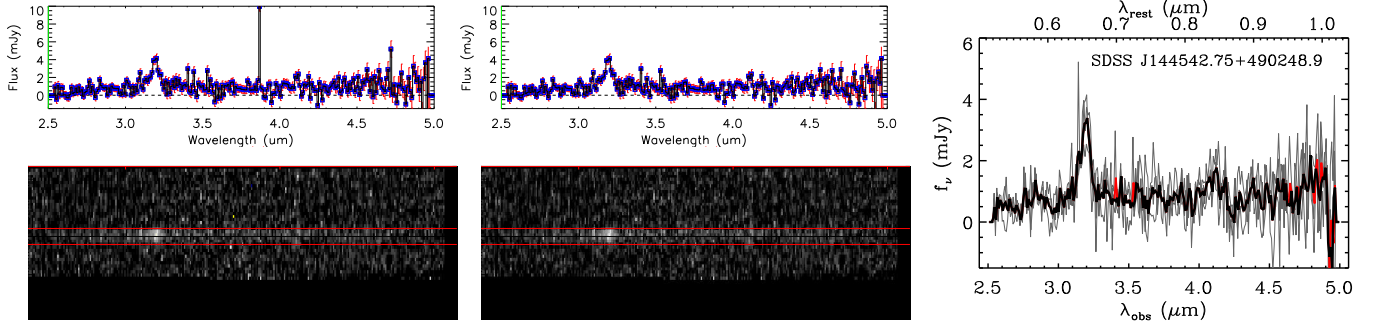


FIG. 2.— A sequential visualization of the data treatment additional to the pipeline processing. From the pipeline processed spectrum (left, 1D at top, and 2D at bottom), hot pixels were rejected (center) before the spectral extraction within the aperture mask (red lines). The multiple pointings extracted (right, gray lines) were stacked (thick black line) with sigma clipping (red line indicates clipped data).

TABLE 1
SUMMARY OF OBSERVATIONS

Program name	Phase	Observed period	Mode	Number of targets	Observed pointings	Rejected pointings
HZQSO	2	Nov 2006–Aug 2007	NG	7	26	2
			NP	5	16	7
QSONG, HQSO2, DPQSO	3	Jun 2008–Jan 2010	NG	147	622	74
			NP	6	11	3

NOTE. — NG and NP stand for NIR grism and prism modes, while a pointing is about 10 minutes long. For the QSONG program data, only the high redshift subsample is noted. The total number of targets is 165 excluding 22 rejected sources from confusion, problems in the spectra or data reduction (section 2.3). 10 objects are either duplicated or different in observing mode or program, yielding an effective total of 155 independent objects.

2007). It offers a low, wavelength dependent spectral resolution (R), where $R=120$ at $3.6\,\mu\text{m}$. This corresponds to a velocity resolution of 2500km s^{-1} in full width at half maximum (FWHM), sampled by a pixel scale of $0.0097\,\mu\text{m}$ in wavelength. The targets were placed in a $1' \times 1'$ rectangular slit aperture to reduce source confusion. The wavelength dependence of R can be expressed as $R=120(\lambda/3.6\,\mu\text{m})$, since the dispersion is nearly a constant (Sakon et al. 2012). Meanwhile, a limited number of NIR prism (NP, $R=19$ at $3.5\,\mu\text{m}$) observations were performed to better catch the fainter continuum and line luminosities. The angular pixel scale is $1.5''$ such that all targets are point-like in our probed redshifts.

The observations were performed under the Astronomical Observation Template (AOT) mode of AOT04, typical for spectroscopic observations. The number of NG pointings per target was normally 3–5 for QSONG, where one pointing observation corresponds to 355 or 400 sec on-source exposure. The number of pointings were determined based on the z -band flux, generally set to be smaller for phase-2 and larger for phase-3 open time programs. The NP observations were performed with usually 1–2 pointings. By the time of termination of the satellite mission, the QSONG program was 85% complete with 144 high redshift quasars observed. The open time programs were complete before the satellite lifetime, adding another 33 targets. Table 1 summarizes the observations.

To supplement the NIR spectra, we compiled the optical spectra of the *AKARI* quasars from the Sloan Digital Sky Survey (SDSS) database (DR10 including both the SDSS-I/SDSS-II and the SDSS-III BOSS data, Ahn et al. 2014), and from observations of APM-UKST quasars (Storrie-Lombardi et al. 1996; Péroux et al. 2001) and Q0000–26 (Schneider et al. 1989), in order

to estimate the C IV line based M_{BH} . Also, we collected broad-band photometric data from optical to mid-infrared (MIR) imaging, for the calculation of the rest-frame UV–optical continuum luminosity of AGNs through SED fitting (section 3.2). The data includes SDSS DR9, 2MASS PSC, UKIDSS DR10, and *WISE* AllWISE releases (Ahn et al. 2012; Skrutskie et al. 2006; Lawrence et al. 2007; Wright et al. 2010), and existing Pan-STARRS, *Spitzer*, and *AKARI* imaging (Hines et al. 2006; Jiang et al. 2006; Oyabu et al. 2009; Jiang et al. 2010, Leipski et al. 2014). The Galactic extinction is corrected for these spectro-photometric data, assuming the total-to-selective extinction ratio of $R_V=3.1$ and using the corrected form (Bonifacio et al. 2000) to the extinction map values of Schlegel et al. (1998). The photometric measurement schemes are different in each survey data, such that host galaxy contamination may not be well subtracted. We keep the diverse magnitude type and imaging resolution of each survey data however, as the high luminosity AGNs yield a compatible set of magnitudes dominated by the central AGN contribution (Jun & Im 2013). The optical spectra and the multi-wavelength imaging data are outlined in Table 2.

2.3. Data reduction

The data were reduced using the automated IDL pipeline package IRC-SPECRED (versions 20110114, 20111121 for phase-2 and 3, Ohyama et al. 2007), where pre-processing (dark, linearity, flat corrections), image registration and coaddition, flux and wavelength calibration, astrometry, spectral extraction, and aperture correction were the main tasks involved. The standard pipeline configuration was adopted, except for the usage of short exposure when the image taken for regis-

tering subframes of spectral data were contaminated by saturated stars. This procedure considerably improved the registration of 2-D spectra in both spatial and wavelength directions. In addition, astrometry of the reference image was upgraded using the 2MASS point source catalog to better extract the faint NIR spectra. Indeed, the zeroth order positions of the spectra and the 2MASS coordinates were visually well aligned for spectral extraction.

Due to the increased number of hot pixels and the background level in the phase-3 data, a non-negligible number of bad pixels remained in the reduced spectral data even after applying the IRC_SPECREP pipeline. To remove the remaining hot pixels, a further data reduction step was taken to obtain a cleaner set of spectra, as depicted in Figure 2. After running the spectroscopic pipeline we subtracted the remaining hot pixels using L.A.Cosmic (van Dokkum 2001) with a threshold of 2.5σ , and combined the 1-D spectra with a 2.5-sigma clipping. This threshold was chosen by visually inspecting the reduced spectra so that the chosen threshold removes the spiky hot pixels efficiently without affecting emission lines. Because the typical FWHM of the broad lines in our AGNs is broader than the spectral resolution of NG (2500 km s^{-1}) but not that of NP ($15,800 \text{ km s}^{-1}$ at $3.5 \mu\text{m}$), the hot pixel rejection and the sigma clipping were applied on the NG data only, while the under-sampled NP emission spectra were kept unchanged. We clipped 2.7% and 2.9% of the NG spectral data through the hot pixel reduction procedure and the combining process respectively.

We extracted the 1-D spectra of 3 pixel width for the phase-2 data to maximize S/N, but the width was widened to 5 pixels for the phase-3 data since the noisier spectra made it difficult to determine the center of the object spectrum. Host galaxy contamination is negligible at the bright luminosities of the sample quasars (S11), enabling flexible extraction widths. Aperture corrections were automatically carried out from the pipeline to derive the total flux, for given respective extraction widths. The pipeline did a fair job of placing the extraction aperture on the right location, but visual inspection showed it necessary to make a -1 pixel shift in spatial direction for 90% of the sample. For 10% of the cases, a different shift of -2 to 1 pixels was necessary. The wavelength zeropoints were determined from the pipeline taking into account the satellite attitude drift and sub-pixel coordinate rounding effects, and we did not apply any further correction as the zeropoints were confined within a 0.5 pixel scatter.

Multiple pointings of the 1-D spectra were stacked for each object, where the NG spectra were interpolated to a fixed wavelength grid, flux averaged, and error rescaled assuming Poisson error statistics. Multiple pointings of NP spectra were stacked without modifying the individual spectrum, due to their poor resolution. We provided secondary flux calibration to the stacked spectra by integrating the *AKARI* fluxes over the *WISE* filter response curves, to match the *W1* and *W2* fluxes together by a constant additive correction. The average and rms scatter of the corrections are $-0.01 \pm 0.19 \text{ mJy}$. Out of 675 pointing observations, we used 589 pointings from 165 objects, since some of the pointing observations were not usable due to contamination of the object spec-

TABLE 2
SUPPLEMENTARY DATA

Name	Wavelength	N	Exposure	Reference
Spectra				
SDSS	3800–9200 Å	111	$\geq 45\text{m}$	1
BOSS	3650–10400 Å	98	$\geq 45\text{m}$	1
APM-UKST	~ 3500 –9000 Å	16	15–60m	2,3
Hale	4500–9000 Å	1	30m	4
Photometry				
SDSS	<i>ugriz</i>	134	54s	5
Pan-STARRS	<i>zy</i>	14	6m	6
2MASS	<i>JHK</i>	76	8s	7
UKIDSS	<i>YJHK</i>	34	40s	8
<i>WISE</i>	<i>W1–W4</i>	161	200s	9
<i>Spitzer</i>	IRAC, MIPS 24 μm	27	17–23m	6,10,11,12
<i>AKARI</i>	IRC 2–11 μm	1	–	13

NOTE. — N is the number of matches to the 165 *AKARI* objects in Table 1. Exposure times are typical values. The references numbered are 1. Ahn et al. 2014; 2. Storrie-Lombardi et al. 1996; 3. Péroux et al. 2001; 4. Schneider et al. 1989; 5. Ahn et al. 2012; 6. Leipski et al. 2014; 7. Skrutskie et al. 2006; 8. Lawrence et al. 2007; 9. Wright et al. 2010; 10. Hines et al. 2006; 11. Jiang et al. 2006; 12. Jiang et al. 2010; 13. Oyabu et al. 2007.

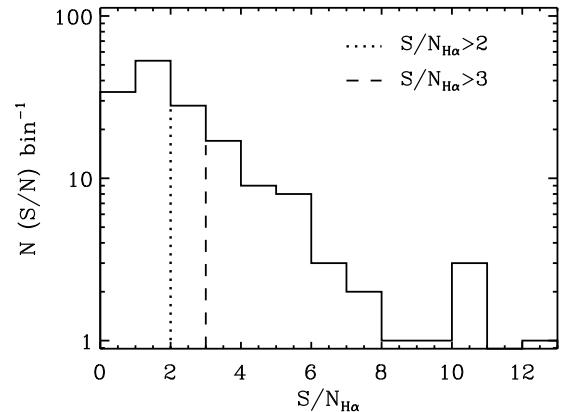


FIG. 3.— The $S/N_{H\alpha}$ distribution of the 160 *AKARI* observed quasars with available S/N (section 3.1). The $S/N_{H\alpha} > 2$ and > 3 cuts are marked in dotted and dashed lines, which are given to limit the measurement of $L_{H\alpha}$ and $\text{FWHM}_{H\alpha}$ respectively (Figure 5).

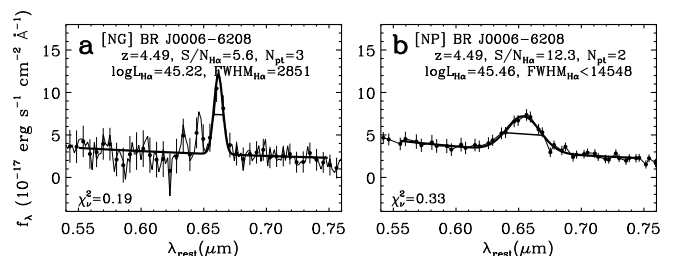


FIG. 4.— Examples of the rest-frame $H\alpha$ emission fitting. On top of the spectra (thin line), the resolution (3 pixel) matched data (dots) and errors, the best-fit to the continuum and $H\alpha$ line emission (thick line), and FWHM are indicated. On the figures, the observation mode, the name of object, the redshift and S/N of the $H\alpha$ emission, the number of *AKARI* pointings, N_{pt} , $H\alpha$ luminosity (erg s^{-1}) and FWHM (km s^{-1}) are printed. The (a) NG and (b) NP observations performed for BR J0006–6208 gives an idea of the enhanced resolution and sensitivity of each spectroscopic mode.

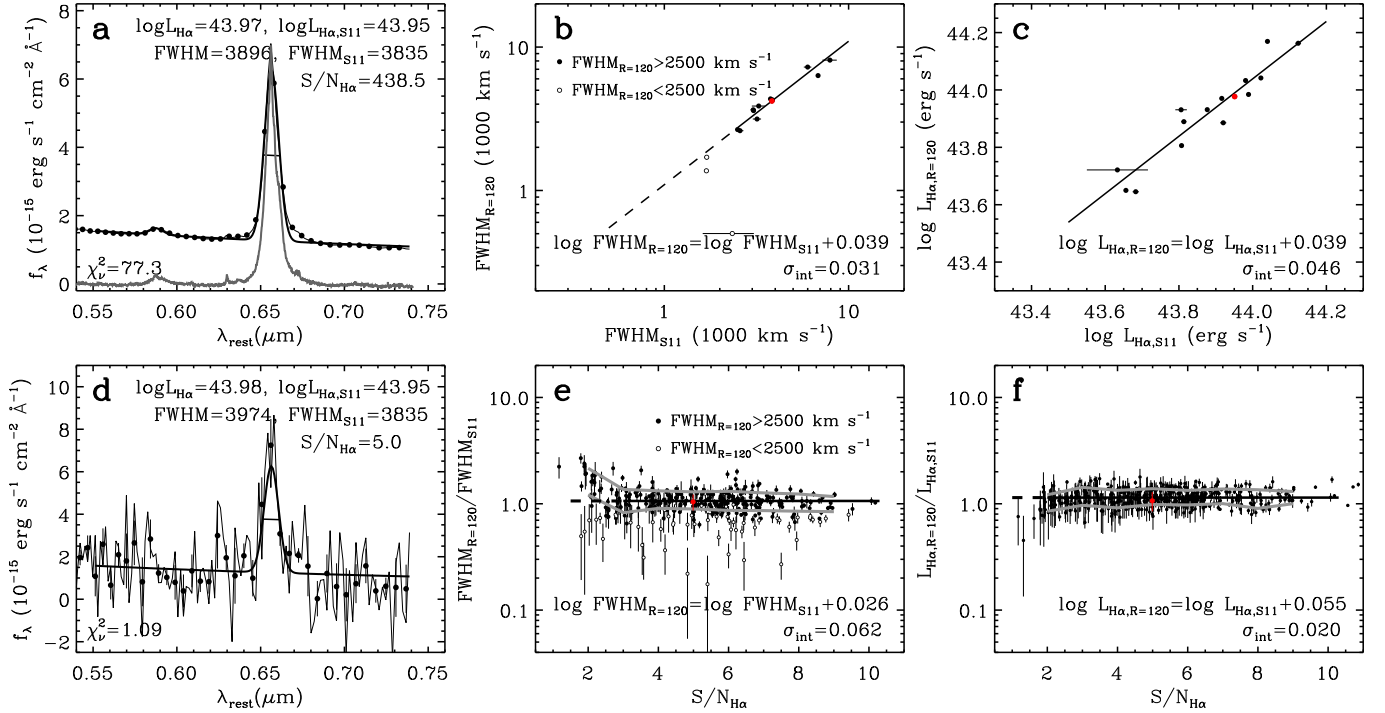


FIG. 5.— Top rows: An example fit to the resolution degraded SDSS spectrum to test the reliability of M_{BH} from *AKARI* observations (left), where the continuum subtracted original SDSS spectrum (gray), and the *AKARI* resolution matched spectrum (black dots) with fit to the simulated data (thick line) are plotted. The broad $\text{FWHM}_{\text{H}\alpha}$ values of $z < 0.3$, $L_{5100} > 10^{45}$ ergs s $^{-1}$ SDSS quasars from S11 are compared to our fitting result of the simulated spectrum (center). Offsets from a 1–1 relation and intrinsic scatter are displayed. For the comparison of FWHMs (center), we divide the sample into $\text{FWHM}_{\text{H}\alpha}$ larger or smaller than 2500 km s $^{-1}$ (filled and open dots) and only use the $\text{FWHM}_{\text{H}\alpha} > 2500$ km s $^{-1}$ data. The red highlighted symbol represents the example on the leftmost panel. Likewise, the $\text{H}\alpha$ luminosity from S11 and our single Gaussian fit to the simulated spectra, are compared (right). Bottom rows: we also test the effect of low S/N to the fitting by adding a set of Gaussian random noise to the resolution degraded spectrum (left). The ratio of our $\text{FWHM}_{\text{H}\alpha}$, $L_{\text{H}\alpha}$ measurements to that from S11 are plotted along the $S/N_{\text{H}\alpha}$ (center and right). For the comparison of FWHMs (center), we remove 49 data points from the plot with $\Delta \text{FWHM}_{\text{H}\alpha} = 0$ usually at $\text{FWHM}_{\text{H}\alpha} < 2500$ km s $^{-1}$. The mean and 1σ offsets are shown in black and gray lines. When calculating the mean and intrinsic scatter of the quantities, only the $S/N_{\text{H}\alpha} > 3$, $\text{FWHM}_{\text{H}\alpha} > 2500$ km s $^{-1}$ data are used for the comparison of $\text{FWHM}_{\text{H}\alpha}$, and the $S/N_{\text{H}\alpha} > 2$ and any $\text{FWHM}_{\text{H}\alpha}$ data for the $L_{\text{H}\alpha}$.

trum by adjacent sources. We excluded frames from the analysis when there was a source that is brighter than the target and its distance from the target is less than the FWHM in spatial direction. In rare occasions, the reduction pipeline did not run properly and such data were not used. In Figure 3, we plot the histogram of $S/N_{\text{H}\alpha}$, the S/N of the $\text{H}\alpha$ emission line measured within $\pm \text{FWHM}_{\text{H}\alpha}$ from the line center.

3. ANALYSIS

3.1. Spectral fitting

We modeled the rest-frame 5500–7500 Å spectra of our AGNs, as a sum of the power-law continuum $f_\lambda = c\lambda^{-(2+\alpha)}$ (where $f_\nu \propto \nu^\alpha$) and the Gaussian $\text{H}\alpha$ emission components. We did not attempt to fit the relatively weak emission features far (He I, [O I]), or near the $\text{H}\alpha$ ([N II] and [S II] doublets, Fe II complex), as they were not detectable under the *AKARI* spectroscopic resolution and sensitivity. Likewise, the $\text{H}\beta$ and [O III] lines were too weak to be detected in most cases and were not fitted accordingly. Examples of the fitting are given in Figure 4, with fitted parameters shown on each panel. We found that 23% of the sample show spiky emission/absorption features around the $\text{H}\alpha$ line from low S/N spectra, which were manually masked out. The $\text{H}\alpha$ line was modeled as a single broad Gaussian with observed FWHMs of 2500–10,000 km s $^{-1}$, for the

inability to clearly resolve the narrow or multiple broad components at $R \sim 120$. The $\text{H}\alpha$ line center was set free within ± 2500 km s $^{-1}$ to the UV line-based redshift from references in Table 2. However, we found four (SDSS J143835.95+431459.2, SDSS J142243.02+441721.4, BR J0307–4945, and SDSS J150654.54+522004.6) exceptions whose $\text{H}\alpha$ line centers were significantly redshifted from the UV-line based redshifts. In these cases, the velocity range was loosened to $\pm 10,000$ km s $^{-1}$, where the $\text{H}\alpha$ showed velocity shifts of 3600–5600 km s $^{-1}$. Next, the measured broad line width FWHM_{obs} , was subtracted by the instrumental resolution FWHM_{ins} (section 2.2) in quadrature, to obtain the intrinsic line width $\text{FWHM} = \sqrt{(\text{FWHM}_{\text{obs}})^2 - (\text{FWHM}_{\text{ins}})^2}$.

Extracting the broad emission line luminosity and width is important in accurately estimating the M_{BH} of AGNs. The limited *AKARI* sensitivity (Figure 3) and spectral resolution could produce systematic bias in the measurement of line parameters such as the line luminosity and width (e.g., Denney et al. 2009). Therefore, we investigated how the low resolution, low S/N spectra systematically affect the results of the spectral fitting. This was done by running Monte Carlo simulations on a set of luminous SDSS quasar spectra (DR7, Schneider et al. 2010) of $R \sim 2000$ to mimic the quality of *AKARI* spectra. Given that the host galaxy contamination to the quasar spectrum is negligible at $L_{5100} > 10^{45}$

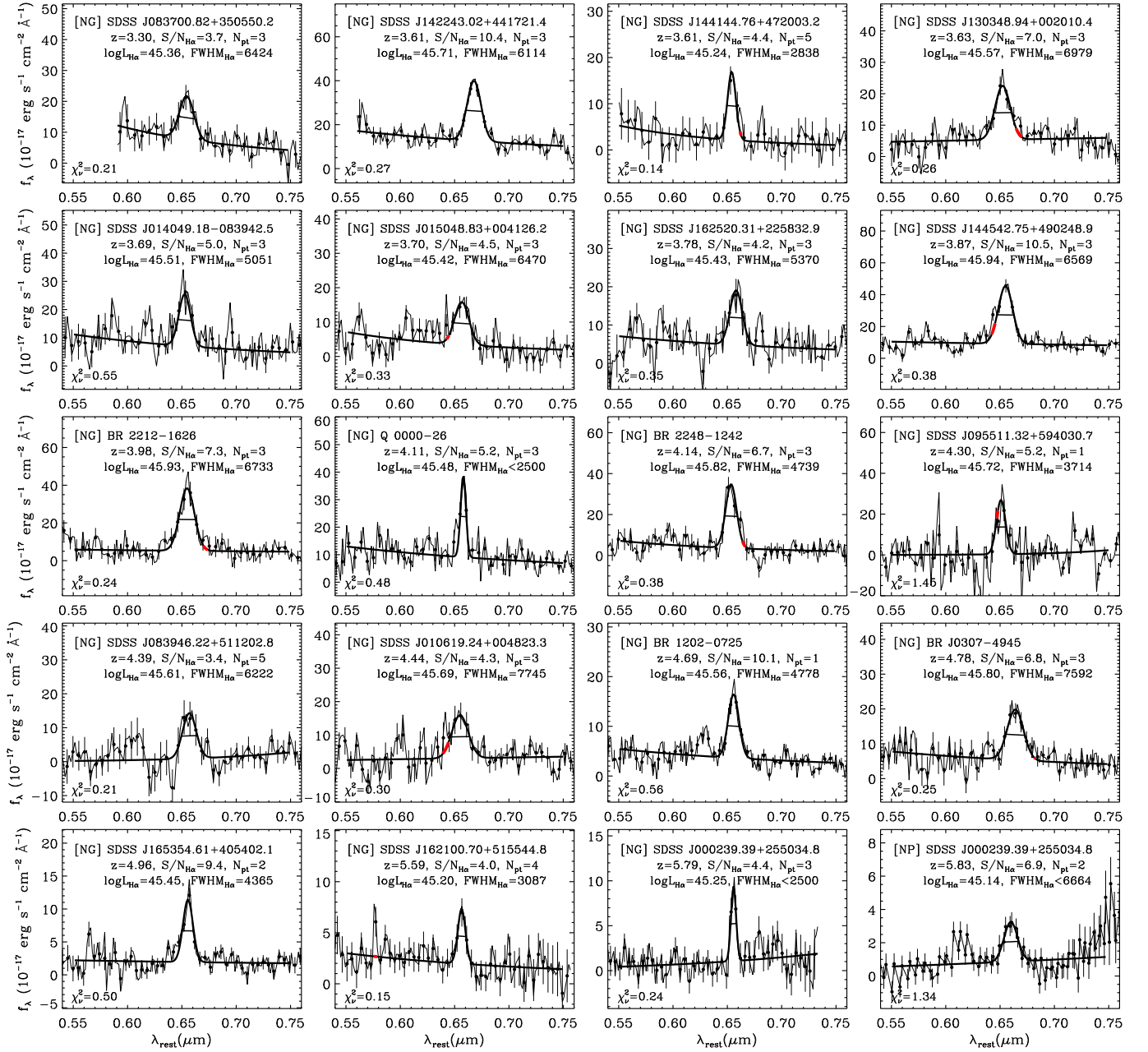


FIG. 6.— Rest-frame H α spectral fitting of selected objects with $S/N_{\text{H}\alpha} > 3$, sorted by the H α redshift. The data point symbols and colors follow the meaning of Figure 4. When there were contaminations to the H α emission, we masked out the region (red).

ergs s^{-1} (Shen et al. 2011, hereafter S11), we collected 15, $L_{5100} > 10^{45} \text{ ergs s}^{-1}$, type-1 quasar spectra with continuum sensitivity of $S/N > 20$, where the H α emission region is present ($z < 0.3$) and well fit ($\chi^2_c < 2$) from S11.

We smoothed the SDSS spectra with a Gaussian function to match the *AKARI* NG resolution, $R=120$, and rebinned the data to match the 3 pixel per resolution sampling of the *AKARI* spectra. First, we looked into the question of how the results get affected with a single Gaussian fit to the emission line in the low resolution of the *AKARI* data. For this, without adding extra noise, we followed the same method to measure the H α line FWHM and luminosity as for the *AKARI* spectra (e.g., Figure 5a). In Figures 5b and 5c we compare the fitted parameters FWHM $_{\text{H}\alpha}$ and $L_{\text{H}\alpha}$ from the smoothed

and binned spectra, to that of the measurement from S11. We find that the FWHM at $> 2500 \text{ km s}^{-1}$ and the line luminosity from the degraded resolution spectra, are remarkably consistent with S11 within ~ 0.04 dex offset and intrinsic scatter σ_{int} , where $\sigma_{\text{int}}^2 = \sum_{i=1}^N [\{y_i - f(x_i)\}^2 - \Delta y_i^2 - \beta^2 \Delta x_i^2] / (N - 1)$ for $f(x) = \alpha + \beta x$ and a set of N data points (x_i, y_i) with measurement errors $(\Delta x_i, \Delta y_i)$. This could bias the M_{BH} measurements up to ~ 0.1 dex in offset and σ_{int} when following the $M_{\text{BH}} \sim L^{0.5} \times \text{FWHM}^2$ behavior, but this is smaller than the typical M_{BH} measurement uncertainty (e.g., Figure 14c). Exceptions to the consistency between the simulated and observed parameters are the line width measurements at $\text{FWHM}_{\text{H}\alpha} < 2500 \text{ km s}^{-1}$, where the simulated FWHM values fall below the extrapolated linear

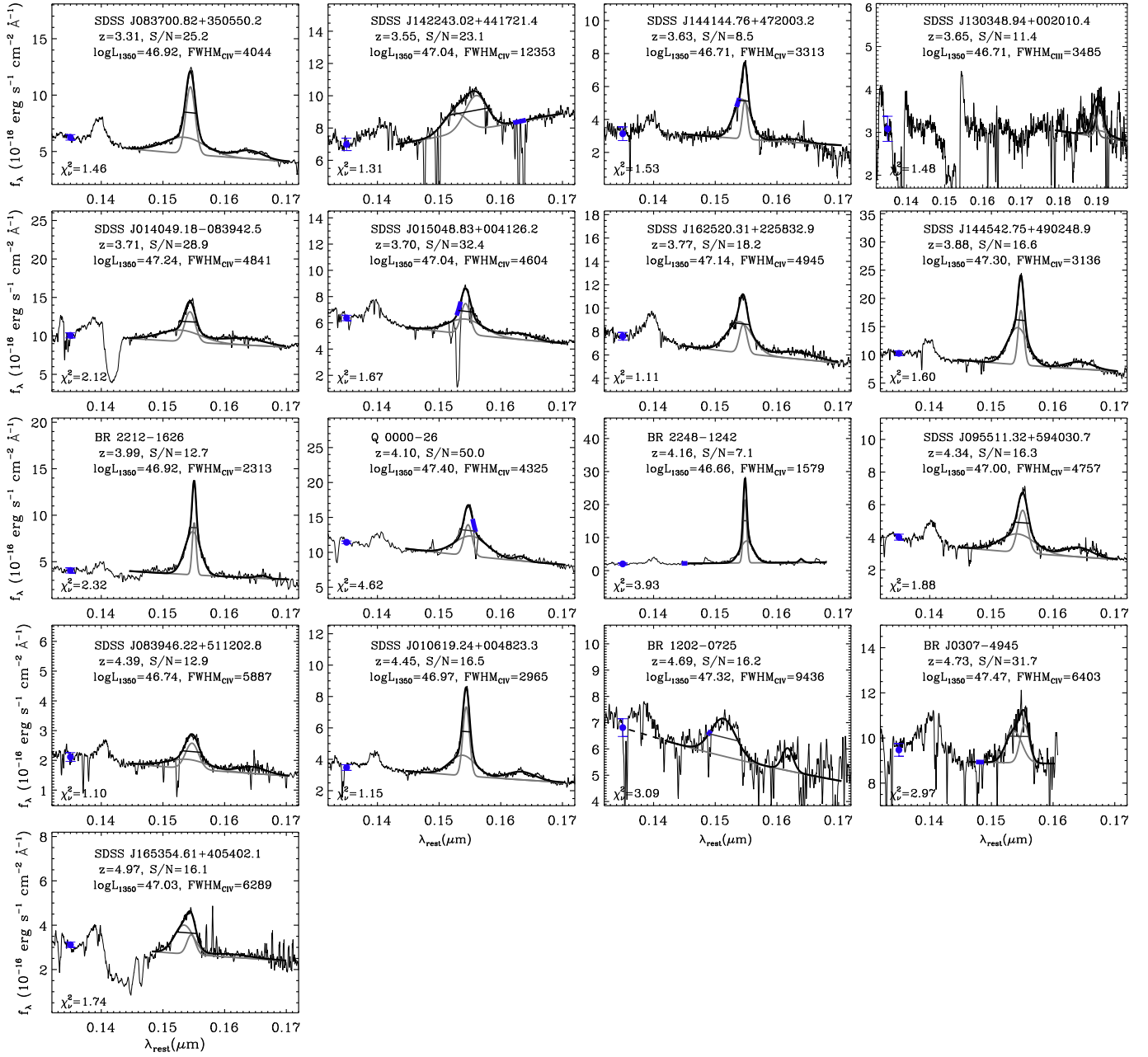


FIG. 7.— Rest-frame C IV spectral fitting of the identical objects in Figure 6 plotted at the same relative location, following the format of the plotted data and printed numbers. Additionally shown are the double Gaussian fit (gray) to the C IV above the continuum, and the 1350 Å monochromatic flux marked on its wavelength (blue dots), with extrapolated values indicated (dashed lines). The spectra are shown when they are available from the listed references in Table 2. When there were contaminations around the C IV emission, we masked out the region (blue). The C IV spectra displayed are smoothed down to $R = 500$, to highlight the spectral features better.

relation. Therefore, we give a $\text{FWHM}_{\text{H}\alpha} > 2500 \text{ km s}^{-1}$ limit to our *AKARI* data to restrict the sample with less biased line width measurements. Meanwhile, the narrow H α and [N II] doublet luminosities of the fiducial SDSS AGNs add up to the broad $L_{\text{H}\alpha}$, by less than 0.01 dex altogether. The weakness of the narrow emission lines in luminous AGNs guarantees that the narrow line contaminations to the degraded resolution spectra are negligible.

Second, to investigate the effect of low S/N to the fitted results, we added a set of random Gaussian noises on top of the degraded resolution spectra. Each SDSS spectrum was repeatedly simulated 30 times adding the random errors, for the $\text{S/N}_{\text{H}\alpha}$ to be distributed down to

the level of *AKARI* S/N. Again, we followed the same method to measure the H α line FWHM (only for the $\text{FWHM}_{\text{H}\alpha} > 2500 \text{ km s}^{-1}$ objects from Figure 5b) and luminosity as for the *AKARI* spectra (e.g., Figure 5d). In Figures 5e and 5f we plot the ratios of $\text{FWHM}_{\text{H}\alpha}$ and $L_{\text{H}\alpha}$ from the degraded resolution and S/N spectra, to that from S11, along $\text{S/N}_{\text{H}\alpha}$. Overall, we find both $\text{FWHM}_{\text{H}\alpha}$ and $L_{\text{H}\alpha}$ to be within 0.02 dex in the mean offset to the values before adding the noises (Figure 5b and 5c), at $\text{FWHM}_{\text{H}\alpha} > 2500 \text{ km s}^{-1}$, $\text{S/N}_{\text{H}\alpha} > 3$ for the $\text{FWHM}_{\text{H}\alpha}$ and $\text{S/N}_{\text{H}\alpha} > 2$ for the $L_{\text{H}\alpha}$. Also, the σ_{int} in $\text{FWHM}_{\text{H}\alpha}$ and $L_{\text{H}\alpha}$ under low S/N are within 0.03 dex to those of the noise-free, effectively unchanging the bias in the M_{BH}

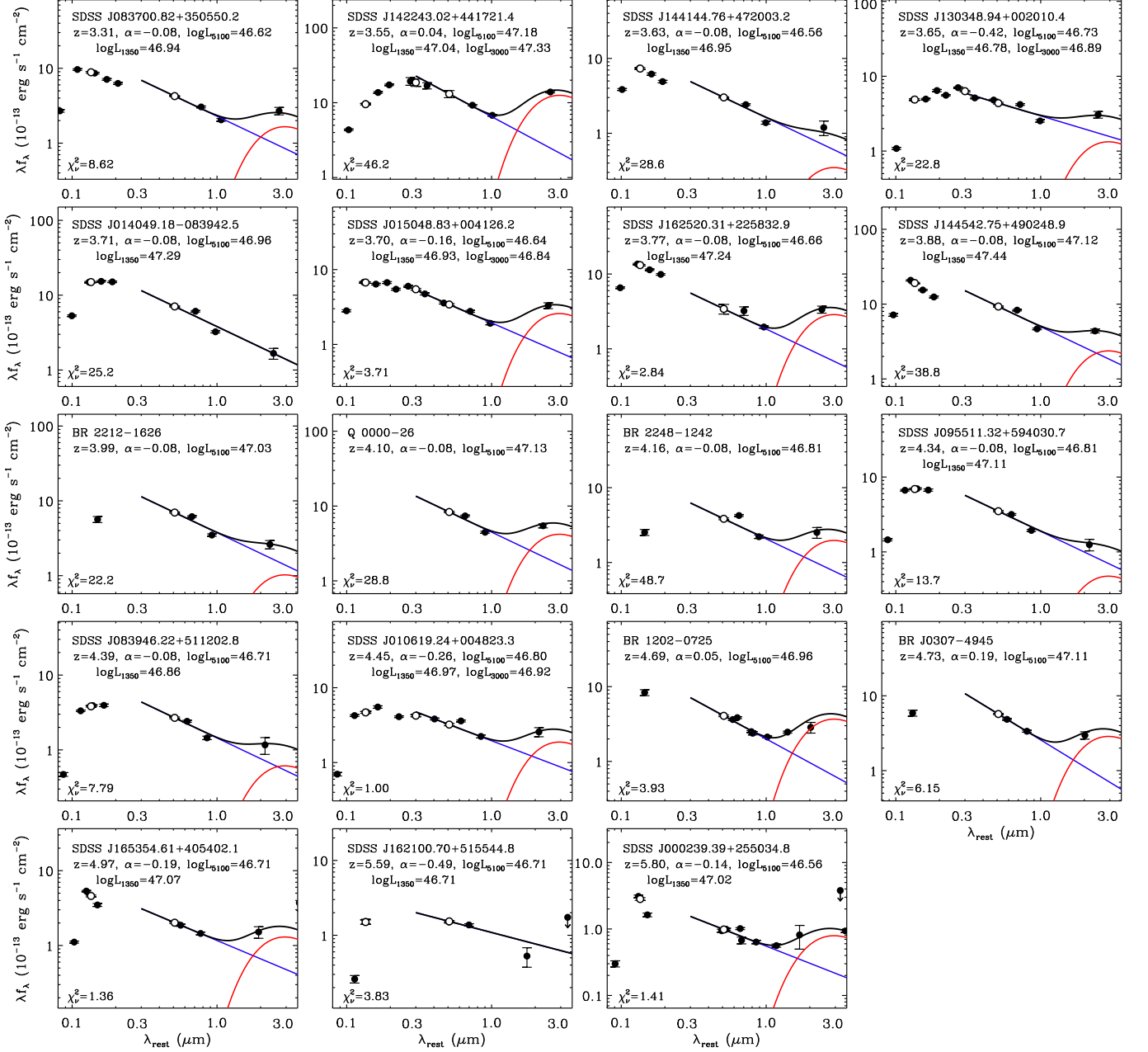


FIG. 8.— Rest-frame UV-NIR broad-band SED of the objects in Figure 6 plotted at the same relative location. The figure shows the observed data points (filled circles) and *WISE* 2σ upper limits (arrows), model fits of the accretion disk (blue line) and the $T=1250$ K dust components (red curve). Also, the monochromatic 1350, 3000, and 5100 Å fluxes are drawn (open circles) when available.

at a similar ~ 0.1 dex. We note that below the S/N or FWHM limit the fitted quantities systematically diverge from those calculated with the original spectrum. Therefore, we conclude that for a given selection of data neither the poor resolution nor sensitivity biases the fitted results by greater than ~ 0.1 dex level of systematic offset or scatter in M_{BH} , and we give the corresponding FWHM and S/N cuts to the *AKARI* data.

Having tested the reliability of fitted quantities under possible systematic biases, we come back to the fitting of the *AKARI* data and find the fit to converge for 160 out of 165 objects. Five failures show noisy spectra near the $\text{H}\alpha$ and were removed. The goodness of the spectral fitting is quantified as the reduced chi-square, and it reaches down to $\chi^2_\nu = 0.33$ in median. The fraction of

AGNs passing the reliability limit for the FWHM measurements ($\text{S/N}_{\text{H}\alpha} > 3$) are 67 % ($N=8$) and 25 % ($N=37$) for the phase-2 and 3 data respectively. Also, we flagged the five NG objects of $\text{FWHM}_{\text{int}} < 2500 \text{ km s}^{-1}$ with an upper limit of 2500 km s^{-1} , and put upper limits on the measured FWHM of the four NP sources (e.g., Figure 4b). In total, there are 43 $\text{FWHM}_{\text{H}\alpha}$ measurements including seven out of nine upper limits, where the excluded two upper limits are NP measurements with overlapping coverage in NG.

Next, we computed the $\text{H}\alpha$ line luminosity $L_{\text{H}\alpha}$ and the 5100 Å continuum luminosity L_{5100} from the spectra, converting the measured rest-frame fluxes using the luminosity distance and assuming isotropic radiation. The $L_{\text{H}\alpha}$ was derived from the Gaussian fit to the observed

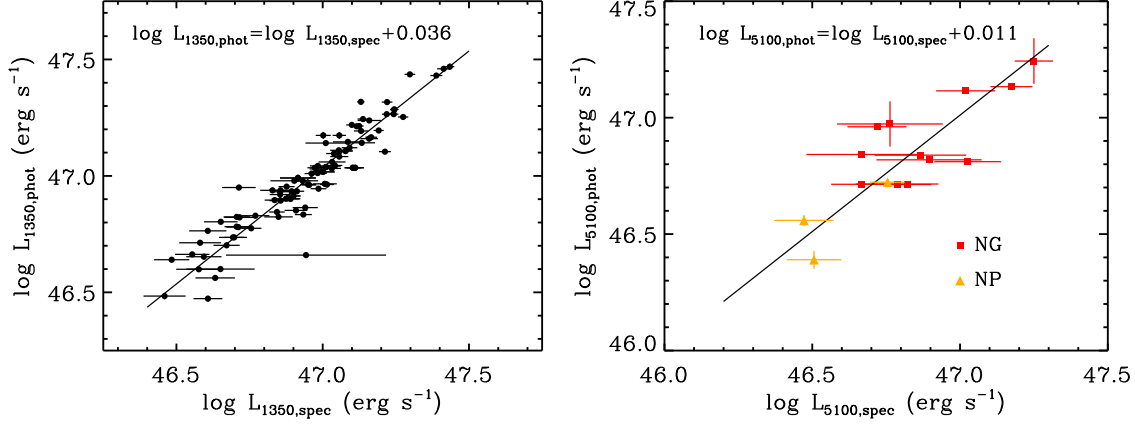


FIG. 9.— Comparison of spectroscopically and photometrically derived monochromatic luminosities, for L_{1350} (left) and L_{5100} (right). The *AKARI* NG/NP data points for the L_{5100} are highlighted with red squares and yellow triangles, respectively. Offsets to the luminosities from a 1–1 relation is displayed.

flux. The reliability limit for $L_{H\alpha}$ ($S/N_{H\alpha} > 2$) is satisfied for 45 % ($N=72$) of the data, while the $S/N_{H\alpha} < 2$ spectra were provided with 2σ upper limits from their given noise levels collected within $\pm 4000 \text{ km s}^{-1}$ of the $H\alpha$ line center. Meanwhile, the L_{5100} was calculated from the average of the rest-frame fluxes at 5000–5200 Å to reduce the measurement uncertainty. The L_{5100} measurements were kept only when $S/N_{5100} > 2$, and the rest of the data were given with 2σ upper limits alike $L_{H\alpha}$. We have less L_{5100} measurements than $L_{H\alpha}$ where only 25 % ($N=41$) meet $S/N_{5100} > 2$, not to mention the limited number of spectra (53 %, $N=88$) covering the rest-frame 5100 Å. Thus, we also derived the L_{5100} alternatively by the photometric SED fitting (section 3.2).

In addition, we fitted the C IV region (rest-frame 1445–1705 Å) of 121 objects with a SED model containing a power law component and double broad Gaussians to model the C IV emission. Also, a single broad Gaussian was used to fit the 1600 Å feature (Laor et al. 1994), and the He II and O III] around 1650 Å altogether since these emission are blended but relatively detached from the C IV. This component was not regarded as a part of the C IV, consistent with the previous studies (e.g., prescription A of Assef et al. 2011; S12). We do not subtract the broad Fe II complex around the C IV emission, as it does not change the FWHM_{CIV} meaningfully (S11). The optical spectra were fitted after carefully masking out the absorption features around the C IV line for 25 % of the spectra. Meanwhile, 16 spectra without error information were fitted assuming the flux error is uniform at all wavelengths, and the rms scatter of the best-fit solution is chosen to be the flux error afterward. Out of five spectra with severe broad absorption line (BAL) features, we fitted the C III] $\lambda 1908$ and used its line width as an effective FWHM_{CIV} surrogate (S12) for two objects, while excluding the remaining three objects from the UV line analysis. In total, we derived FWHM_{CIV} and L_{1350} for 118 objects. The L_{1350} and its error were calculated from the average of the rest-frame 1350 ± 15 Å fluxes to avoid contamination from narrow absorption, while for seven BAL quasars we extrapolated the continuum around the C IV emission to 1350 Å. When spectra of an object were available from both SDSS-I/SDSS-II and BOSS, we performed the fit to the spectra from each dataset sepa-

ately, and took the average of the parameter values from the independent fits. We plot examples of the spectral fitting of the $H\alpha$ region in Figure 6, and of the C IV in Figure 7.

3.2. Broad-band SED fitting

The photometry datasets in Table 2 cover a wide wavelength range in broad-band filters from u -band through $24 \mu\text{m}$, thus we fitted the broad-band SEDs to provide further information on the AGN continuum luminosities. Under a photometric sensitivity limit of $S/N > 5$ for the observed optical–NIR and $S/N > 2$ in the MIR, and further rejecting the 2MASS data with a single filter detection, we modeled the SED in the rest-frame $0.3\text{--}5 \mu\text{m}$ including 5 data points on average, as a sum of a power law continuum and a black body emission from hot dust of $T=1250 \text{ K}$ (e.g., Jun & Im 2013). For 53% of the case ($N=87$) for which there were no rest-frame $0.3\text{--}0.6 \mu\text{m}$ data points available, we used the average continuum slope $\alpha = -0.08$ of luminous SDSS quasars (Jun & Im 2013). The uncertainty in fixing the continuum slope was tested from objects that cover the rest-frame $0.3\text{--}0.6 \mu\text{m}$, and comparing the L_{5100} with and without fixing the α value. The test yields the L_{5100} to be offset by $-0.01 \pm 0.04 \text{ dex}$ when α is fixed, compared to when α is free. Since the offsets are small, we find our method to fix the α when missing the photometric coverage near the rest-5100 Å, to be reliable in tracing the L_{5100} . In addition, considering the filter bandwidths and the AGN line equivalent widths from Vanden Berk et al. (2001), we find the $H\alpha$ to be the only line that meaningfully contributes to the rest-optical photometry over the continuum emission (by $> 0.03 \text{ dex}$). Thus, we removed the data point enclosing the $H\alpha$ emission when the χ^2_ν containing that data point became larger than that without. Through this procedure we obtained 164 photometrically derived L_{5100} , while removing one object without any detections in the rest-frame $0.3\text{--}5 \mu\text{m}$ under our sensitivity limit. Examples of the broad-band SED fitting are shown in Figure 8. The reduced chi-square values have a median of $\chi^2_\nu = 3.6$, is acceptable given the simplification of the SED model that does not take into account the emission line features, and the general agreement of the fit to the data demonstrated in Figure 8.

Meanwhile, we interpolated the broad-band SED

TABLE 3
CONTINUUM AND LINE BASED PROPERTIES OF THE SAMPLE

Name (1)	z_{ref} (2)	$z_{\text{H}\alpha}$ (3)	$\log L_{1350}$ (4)	$\log L_{5100}$ (5)	$\log L_{\text{H}\alpha}$ (6)	$\text{FWHM}_{3, \text{C IV}}$ (7)	$\text{FWHM}_{3, \text{H}\alpha}$ (8)	$\log M_{\text{BH}, \text{C IV}}$ (9)	$\log M_{\text{BH}, \text{H}\alpha}$ (10)
SDSS J000239.39+255034.8	5.80	5.79	99.00 ± 99.00	46.56 ± 0.02	45.14 ± 0.11	99.00 ± 99.00	2.50 ± -1.00	99.00 ± 99.00	9.25 ± -1.00
Q 0000-26	4.10	4.11	47.40 ± 0.01	47.13 ± 0.01	45.48 ± 0.13	4.33 ± 0.28	2.50 ± -1.00	9.88 ± 0.22	9.56 ± -1.00
SDSS J000552.34-000655.8	5.85	5.85	99.00 ± 99.00	46.04 ± 0.08	44.75 ± 0.32	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR J0006-6208	4.45	4.49	46.90 ± 0.02	46.71 ± 0.00	45.46 ± 0.06	11.33 ± 1.33	2.85 ± 0.85	10.48 ± 0.24	9.46 ± 0.31
SDSS J001115.23+144601.8	4.97	99.00	99.00 ± 99.00	46.78 ± 0.02	47.97 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00

NOTE. — Catalog of the properties derived for the *AKARI* quasars, sorted by right ascension. Column 1: Target name; Column 2: Redshift from references; Column 3: Redshift measured from $\text{H}\alpha$; Column 4: 1350 Å luminosity and its uncertainty; Column 5: 5100 Å luminosity and its uncertainty; Column 6: $\text{H}\alpha$ luminosity and its uncertainty; Column 7: FWHM of the C IV line and its uncertainty; Column 8: FWHM of the $\text{H}\alpha$ line and its uncertainty; Column 9: M_{BH} from the C IV line and its uncertainty; Column 10: M_{BH} from the $\text{H}\alpha$ line and its uncertainty. The units for L , FWHM, and M_{BH} are ergs s^{-1} , 1000 km s^{-1} , and M_{\odot} . Columns 9 and 10 are from Equations (10) and (7), respectively. Empty parameters are entered as 99 and upper limits are given with errors of -1.
(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

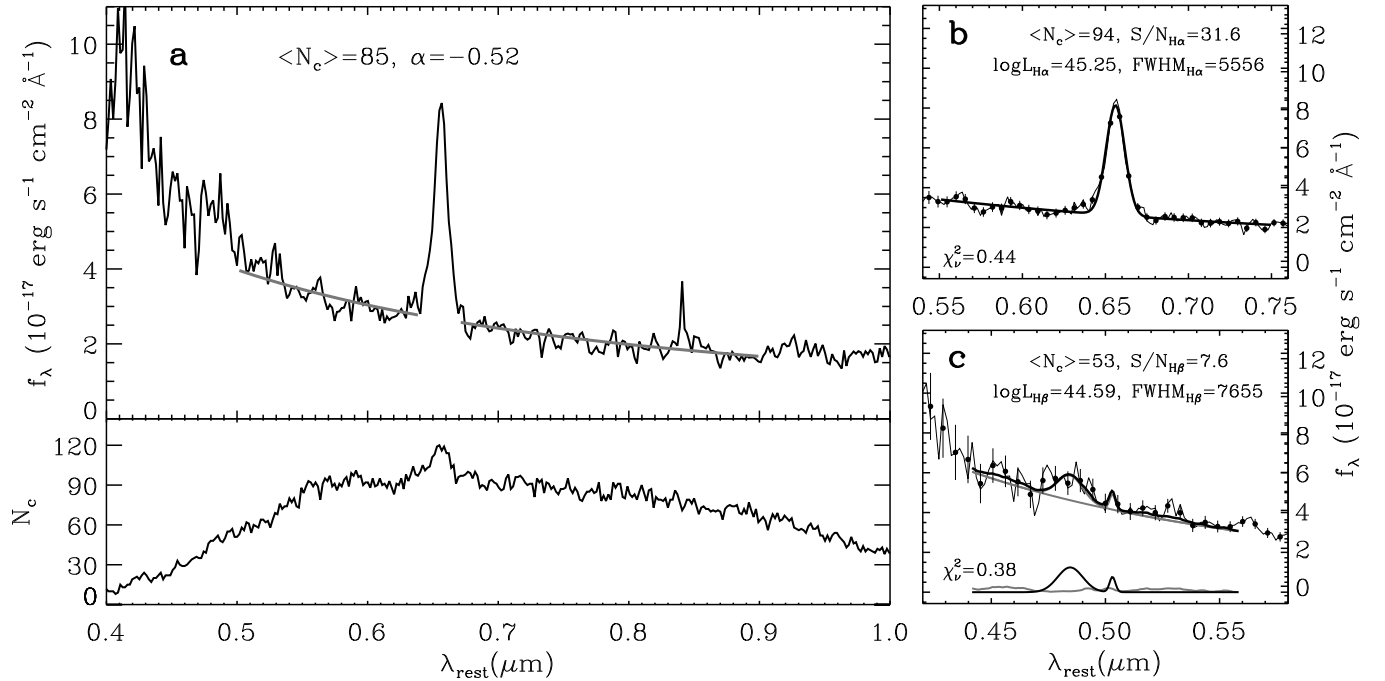


FIG. 10.— Composite spectra from *AKARI* observations. (a) The spectra within 0.4–1 μm (top), and the number of spectra used to construct the composite, N_c , plotted against wavelength (bottom). The mean N_c and the continuum slope from the 0.5–0.9 μm region are printed. (b) $\text{H}\alpha$ region fit of the composite spectra. (c) $\text{H}\beta$ region fit of the composite spectra. The sum of the continuum, $\text{H}\beta$ emission, and the Fe II complex (thick black line), and the continuum (gray line) are overplotted on the data, while the $\text{H}\beta$ /[O III] and the Fe II emission components are separately plotted below the spectrum in black and gray lines.

around the rest-frame 1350 and 3000 Å to obtain L_{1350} and L_{3000} . For this, we used objects with more than two data points in the rest-frame 500–2500 Å for L_{1350} or 2000–6000 Å for L_{3000} . The interpolation is done linearly to the data points, and we obtain 137 L_{1350} and 47 L_{3000} values. The continuum luminosities derived by photometric and spectroscopic methods roughly agree with each other as shown in Figure 9 for L_{1350} and L_{5100} , though there are not enough data points ($N=2$) to plot for L_{3000} . Likewise to L_{5100} , we calculated the level of C IV or Mg II line contamination to the L_{1350} and L_{3000} from broad-band photometry. The C IV and Mg II elevates the observed broad-band flux by up to 0.06 and 0.04 dex. Thus, it is possible that the photometry embracing the broad UV emission is boosted by more than the typical measurement error in L_{1350} and L_{3000} , which are 0.02 and 0.03 dex, respectively. Between the spec-

troscopically and photometrically derived continuum luminosities, we will use in the following discussion the spectroscopically derived L_{1350} and the photometrically derived L_{3000} and L_{5100} . We do so because the line contaminations near 1350 Å through the broad-band photometry can be significantly larger than the measurement uncertainty, while it is not so around the Mg II and $\text{H}\alpha$ lines. The large error in L_{5100} from *AKARI* spectra is also another reason why we opt to use L_{5100} from the broad-band SED fitting. The fitted properties from this section, and the M_{BH} to appear in section 4, are listed in Table 3. For objects with both NG and NP observations, we use the $z_{\text{H}\alpha}$ and $\text{FWHM}_{\text{H}\alpha}$ from the NG and the $L_{\text{H}\alpha}$ from the NP when $S/N_{\text{H}\alpha} > 3$, while we list only the values from the NG otherwise.

4. RESULTS

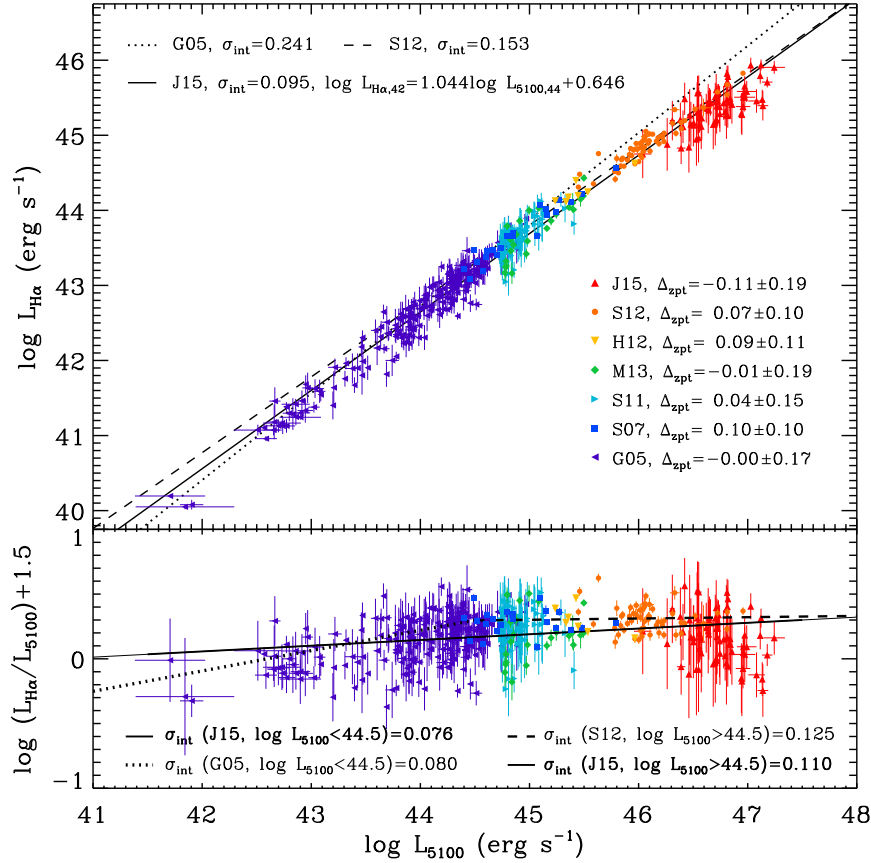


FIG. 11.— The L_{5100} – $L_{H\alpha}$ relation of AGNs (top), and its projection on the L_{5100} – $L_{H\alpha}/L_{5100}$ (bottom), from combined references. The references abbreviated on the plot are summarized in Table 4, where all measurements from references are converted to our adopted cosmology. We limited the G05 data to $L_{5100} < 10^{44.73}$ ergs s $^{-1}$ to avoid overlap with the S11 data. The L_{5100} – $L_{H\alpha}$ relation from G05, S12, and J15 are shown in dotted, dashed, and solid lines respectively. The zeropoint offset and rms scatter of each literature data with respect to our relation are denoted as Δ_{zpt} . The intrinsic scatter (σ_{int} , dex) of the entire data with respect to the G05, S12, and J15 relations (top), and of the data divided by $L_{5100} = 10^{44.5}$ ergs s $^{-1}$ (bottom), are shown. The L_{5100} in M13 and J15 are from photometric SED fitting (section 3.2), while the rest are spectroscopically derived from each reference. The $L_{H\alpha}$ are from broad emission for all references but for the G05, S07, and S11 data, where the narrow component is included.

4.1. Composite spectra

To investigate the overall rest-optical spectral properties of the sample, we construct the composite *AKARI* spectra. Out of 154 objects observed with NG, 127 are used for the composite construction after removing 27 spectra due to a mild level of confusion from neighbor source spectra, negative continuum levels, or strong fluctuations near the $H\alpha$ due to warm pixels. The composite is constructed by normalizing the spectra at 5100 Å and taking their error weighted mean to maximize the S/N. Each spectral flux and error were de-redshifted, and rebinned to a common wavelength scale of 18 Å per bin which is equal to that of the *AKARI* at rest-frame $H\alpha$. In Figures 10a–10c, we plot the composite spectrum, and zoomed-in fit to the $H\alpha$ and $H\beta$ regions, respectively. The $H\alpha$ emission is prominent in the composite spectrum, and the spiky feature at 8400 Å is an artifact from a single spectrum with high S/N.

We determine the continuum slope from the 0.5–0.9 μ m region from Figure 10a, where the number of spectra used to construct the composite exceeds 60. The slope $\alpha = -0.52 \pm 0.06$ (where $f_\nu \propto \nu^\alpha$) is close to $\alpha = -(0.37-0.48)$ of Glikman et al. (2006) determined through the composite of local luminous quasars at similar wavelengths, indicating a similarity in the rest-optical con-

tinuum shape of luminous type-1 quasars with respect to redshift. Interestingly, we detect a sign of the $H\beta$ emission from Figure 10c. The $H\beta$ region was fitted with the Boroson & Green (1992) Fe II template, following the method of Shen et al. (2008). We find $L_{H\alpha}/L_{H\beta} = 4.5 \pm 1.6$, which is roughly consistent with 3.6 ± 1.4 from luminous $z \sim 2$ quasars (S12) or the model broad Balmer line decrement of AGNs at $T=10,000$ – $12,000$ K, $L_{H\alpha}/L_{H\beta} = 3.6$ – 8.8 (Osterbrock 1989). Since the Balmer decrement value and the $S/N_{H\alpha}$ of each *AKARI* spectrum suggest that the strongest $H\beta$ in our individual spectrum would appear as $S/N_{H\beta}=1$ – 2 , we do not expect the $H\beta$ emission to be individually detected, consistent with the visual inspection in section 3.1. Apart from this, we do not find hint of other emission lines in the composite spectra.

4.2. Luminosity and Line Width Scaling Relations

The derivation of continuum and line luminosities for distant, luminous quasars allows us to examine the universality of the luminosity scaling relations over a wide range of redshifts and luminosities. Starting from the L_{5100} – $L_{H\alpha}$ relation, we plot in Figure 11 our derived data points and the data taken from literatures (Greene & Ho 2005, hereafter G05; Shang et al. 2007, hereafter S07; S11; Ho et al. 2012; Matsuoka et al. 2013;

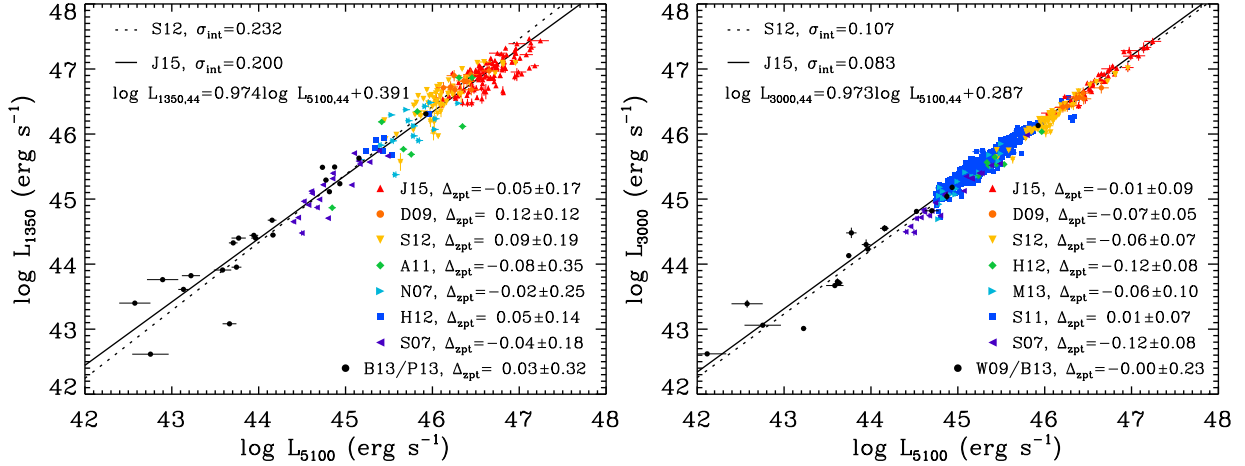


FIG. 12.— The L_{5100} – L_{1350} and L_{5100} – L_{3000} relations of AGNs from combined references. The references abbreviated on the plot are summarized in Table 4, where all measurements from references are converted to our adopted cosmology. The L_{1350} of the N07 sample are searched from S11. The L_{5100} , L_{3000} , and L_{1350} are derived from spectra, except for the L_{5100} , L_{3000} from J15 where they are from photometric SED fitting (section 3.2). We assign a modest 10% error for L_{5100} of the N07 data and L_{1350} , L_{5100} of the A11 data, and 20% error for the L_{1350} , L_{3000} , and L_{5100} from D09 data, from visual inspection of their spectra. We removed the highly variable object 3C 390.3 from the B13 data, and additional two objects in S07 data that overlap with B13. The L_{5100} – L_{1350} , L_{5100} – L_{3000} relations from S12 and J15 are shown in dotted and solid lines respectively. The zeropoint offset and rms scatter of each literature data with respect to our relations, are indicated.

TABLE 4
DYNAMIC RANGE OF REFERENCES

Reference	z	$\log L_{5100} \text{ (ergs s}^{-1}\text{)}$
G05	<0.35	41.7–45.0
N07	2.3–3.5	45.2–46.3
S07	0.08–0.40	44.4–45.8
M08	0.35–0.37	43.6–44.4
D09	1.1–2.2	46.1–46.7
W09	0.00–0.16	42.0–45.9
A11	1.4–3.6	44.8–46.5
S11	0.08–0.80	43.8–46.4
H12	1.4–1.5	45.2–46.0
S12	1.5–2.2	45.4–47.0
B13	0.00–0.29	42.0–45.9
M13	0.7–1.7	44.8–45.5
P13	0.01–0.23	42.6–45.9
J15	3.3–6.2	46.0–47.2

NOTE. — The abbreviated references are Greene & Ho 2005 (G05); Netzer et al. 2007 (N07); Shang et al. 2007 (S07); McGill et al. 2008 (M08); Dietrich et al. 2009 (D09); Wang et al. 2009 (W09); Assef et al. 2011 (A11); Shen et al. 2011 (S11); Ho et al. 2012 (H12); Shen & Liu 2012 (S12); Bentz et al. 2013 (B13); Matsuoka et al. 2013 (M13); Park et al. 2013 (P13); and this work (J15).

S12) that cover a range of L_{5100} and z , as summarized in Table 4. Our *AKARI* data extends the relation at the high redshift ($z > 3.3$) and high luminosity end ($L_{5100} > 10^{46} \text{ ergs s}^{-1}$). To minimize the host galaxy contribution to the AGN luminosities, we chose AGNs with host contamination $< 20\%$ in L_{5100} , $L_{H\alpha}$ for some datasets (G05; S07), while we plotted only the $L_{5100} > 10^{44.73} \text{ ergs s}^{-1}$ data for the rest of references that meet $< 10\%$ in host contamination (S11). Meanwhile, the broad $L_{H\alpha}$ could contain the narrow component for *AKARI* data, while the broad and narrow line luminosities are combined for the G05, S07, and S11 data too. We find that the contribution from the narrow component to $L_{H\alpha}$ estimated from section 3.1 and the references, is negligible (2% and $< 10\%$, respectively), allowing us to consider $L_{H\alpha}$ to be approximately the line luminosity of the broad line component.

Figure 11 shows a remarkable correlation between L_{5100} and $L_{H\alpha}$ even when AGNs are drawn from various samples covering a wide range of redshifts and luminosities. This strongly suggests that the physics governing the correlation is the same for low and high luminosity AGNs, and there is no strong evolution in the relation from $z = 0$ to $z = 6$ over the range explored here. The *AKARI* data points are mildly below the relation where the offset could indicate a growing population of weak emission line quasars at high redshift (Fan et al. 1999; Diamond-Stanic et al. 2009), but the overall deviations are within the scatter of the data. The possible downturn in the correlation at very high luminosity ($L_{5100} \sim 10^{47} \text{ ergs s}^{-1}$) produces a ~ 0.2 dex offset in $\log (L_{H\alpha}/L_{5100})$, but this downturn affects the M_{BH} estimates only by ~ 0.1 dex. The possible downturn could be caused by cold accretion disks of slowly spinning, extremely high mass BHs (Laor & Davis 2011; Wang et al. 2014), but we will leave the investigation of the possible downturn as a subject of a future work as its effect on M_{BH} estimates is small. We fitted the relation using the linear regression with bivariate correlated errors and intrinsic scatter (BCES, Akritas & Bershady 1996)¹⁵, to find the following result,¹⁶

$$\log L_{H\alpha,42} = (1.044 \pm 0.008) \log L_{5100,44} + (0.646 \pm 0.011). \quad (2)$$

The best-fit relation fits the entire data with $\sigma_{\text{int}} = 0.095$ dex. The flux–flux relation of 5100 Å continuum and $H\alpha$ shows almost identical slope and the intrinsic scatter, suggesting that the tight correlation in Equation (2) is not due to a sample selection effect. The data points at three redshift intervals, namely $0 < z < 0.8$ (G05, S07, and S11) for $41.7 < \log L_{5100} < 46.4$,

¹⁵ Throughout this paper we use the BCES fit to derive the linear relations.

¹⁶ Throughout this paper we use subscript numbers to the luminosity to indicate its wavelength and unit, such as $L_{5100,44} = L_{5100} \text{ Å}/10^{44} \text{ ergs s}^{-1}$.

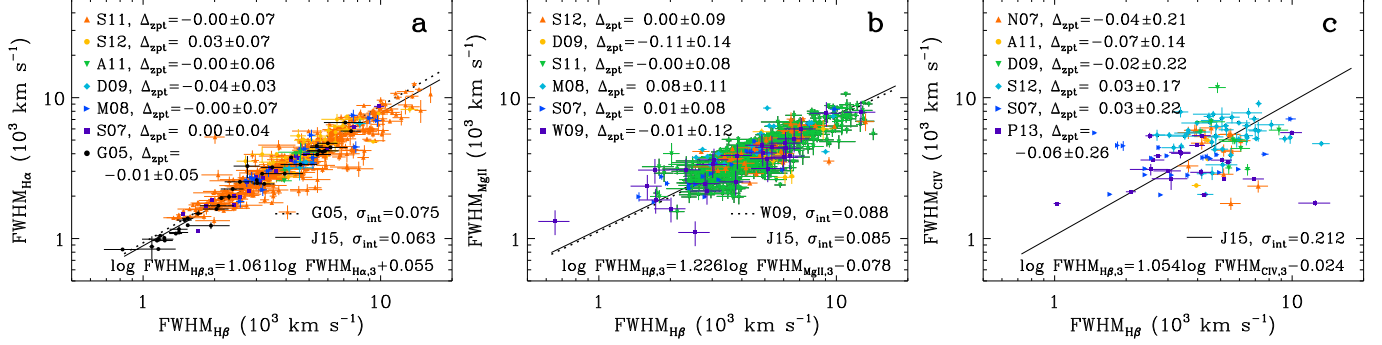


FIG. 13.— The $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{H}\alpha}$, $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{MgII}}$, and $\text{FWHM}_{\text{H}\beta}$ – FWHM_{CIV} relations of AGNs from combined references. The references abbreviated on the plot are summarized in Table 4. The *AKARI* data are not present since the sample do not have simultaneous coverage of the FWHMs plotted. We limited the G05 data to $L_{5100} < 10^{43.83} \text{ ergs s}^{-1}$ to avoid overlap with the S11 data. We assign a 10% error for $\text{FWHM}_{\text{H}\beta}$, FWHM_{CIV} of the N07 data, following their argument. We removed one object in S07 data that overlaps with W09, and two objects in S07 that overlap with P13. The $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{H}\alpha}$ relation of G05, and $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{MgII}}$ relation of W09 are shown in dotted lines, while the relations from this work are shown as solid lines. The zeropoint offset and rms scatter of each literature data with respect to our relations, are indicated. The FWHMs are of the broad emission line.

$0.7 < z < 2.2$ for $44.8 < \log L_{5100} < 47.0$ (H12, S12, and M13), and $3.3 < z < 6.2$ for $46.0 < \log L_{5100} < 47.2$ (J15), overlap with each other and show no evolution. This suggests that Equation (2) is universal, and not due to a distance effect like the Malmquist bias.

To examine the universality of the relation further, we discuss how the L_{5100} – $L_{\text{H}\alpha}$ relation of G05 at $z \sim 0$ of lower luminosity AGNs and S12 at $z \sim 2$ of higher luminosity AGNs fare with each literature values. First, the G05 relation can describe the L_{5100} – $L_{\text{H}\alpha}$ relation of S07, S11, and M13 AGNs over the overlapping luminosity interval ($L_{5100} \lesssim 10^{45} \text{ ergs s}^{-1}$). When extrapolated to higher luminosity, it starts to deviate from the data points regardless of redshift. Likewise, the S12 relation can describe the L_{5100} – $L_{\text{H}\alpha}$ relation down to $L_{5100} \lesssim 10^{44.5} \text{ ergs s}^{-1}$ including the S07 data at $z < 0.4$ and our AGNs at $z > 3.3$. However, when applied to the entire datasets, the G05 and S12 relations show deviations from the data at high and low luminosity regions, respectively. Consequently, both relations produce $\sigma_{\text{int}} = 0.15\text{--}0.24 \text{ dex}$ against the data which is worse than 0.095 dex of our L_{5100} – $L_{\text{H}\alpha}$ relation.

To check if the inconsistency in the G05 and S12 relations at the faint and luminous end arises from a possible break in the relation itself, we considered the case where the slope changes at $L_{5100} \sim 10^{44.5} \text{ ergs s}^{-1}$, where the G05 and S12 relations meet. For this, we computed the σ_{int} of the $L_{5100} < 10^{44.5} \text{ ergs s}^{-1}$ and $L_{5100} > 10^{44.5} \text{ ergs s}^{-1}$ data against the G05, S12, and our relation in Equation (2), and examined if our simple relation is any worse than the combination of G05 and S12 relations with a break at $L_{5100} = 10^{44.5} \text{ ergs s}^{-1}$. As indicated in the lower panel of Figure 11, the σ_{int} values against our relation is comparable to or slightly smaller than the σ_{int} against the G05 or S12 relations at low and high luminosities respectively. This suggests that there is no strong need for a broken power-law form of the L_{5100} – $L_{\text{H}\alpha}$ relation, and a simple relation of Equation (2) can be employed to describe the response of broad line region to the incident continuum emission over the covered redshift and luminosity ranges in Figure 11 and Table 4.

Likewise, we plot in Figure 12 the L_{5100} – L_{1350} and L_{5100} – L_{3000} relations from the luminosities derived in section 3 and taken from references. Apart from the lit-

erature data where the host contamination in L_{5100} is estimated to be $< 20\%$ (S07), or minimized by Hubble Space Telescope (HST) observations, we limit the literature sample with $L_{5100} > 10^{44.73} \text{ ergs s}^{-1}$ to keep the host galaxy contamination below 10% (S11). Meanwhile, we replaced the L_{5100} of Wang et al. (2009) and Park et al. (2013) with the HST data from Bentz et al. (2013), while only including the AGNs with $< 20\%$ host contamination in L_{5100} . Similarly to the L_{5100} – $L_{\text{H}\alpha}$ relation, we do not find any evolution in the L_{5100} – L_{1350} and L_{5100} – L_{3000} relations for a particular set of data, and we find the best-fit correlation to be,

$$\begin{aligned} \log L_{1350,44} &= (0.974 \pm 0.023) \log L_{5100,44} \\ &\quad + (0.391 \pm 0.053) \\ \log L_{3000,44} &= (0.973 \pm 0.010) \log L_{5100,44} \\ &\quad + (0.287 \pm 0.013). \end{aligned} \quad (3)$$

The σ_{int} values (dex) to these best-fit relation, as well as the relation with respect to S12 are presented in Figure 12. Like for the $L_{\text{H}\alpha}$ – L_{5100} relation, the flux–flux relations show virtually identical slopes and intrinsic scatter to the luminosity–luminosity relations, showing again that sample selection effect is not a main driver for the relations in Equation (3).

Finally, we compare the broad line FWHM of $\text{H}\beta$, $\text{H}\alpha$, Mg II , and C IV in Figure 13 in order to calibrate the M_{BH} from multiple line based recipes and to check for any evolution in the FWHM relations. For good comparison of FWHMs, we restricted the mixed samples to have the fractional errors of FWHM less than 20%, while additionally limiting the S/N and reduced chi-square of the SDSS spectra to be $\text{S/N} > 20$ and $\chi^2_{\nu} < 2$. In Figure 13a, we fit the $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{H}\alpha}$ relation from the collected data, where we find the offset of each reference data to this relation to fall within each scatter. The σ_{int} of all the data to our relation, 0.063 dex , is smaller than when using the relation from G05.

Second, we derive the $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{MgII}}$ relation. Since it is debatable whether to subtract the narrow component for the Mg II line width measurement (e.g., S11), we followed Jun & Im (2013) to average the $\text{FWHM}_{\text{MgII}}$ derived with and without the subtraction of the narrow component. Also, the slope and constant of the relation

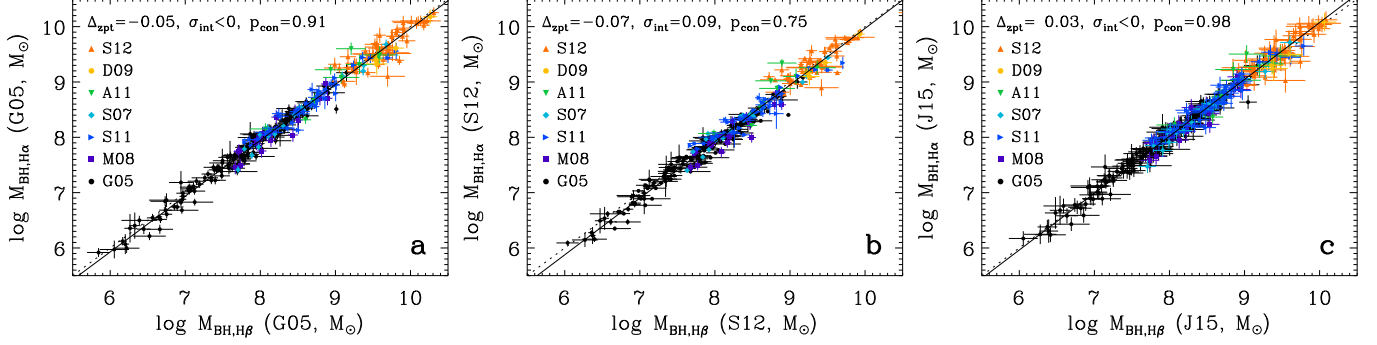


FIG. 14.— The comparison of Balmer M_{BH} relations of AGNs, from combined references. The references abbreviated on the plot are summarized in Table 4. All measurements are converted to our adopted cosmology and f -factor. We limited the G05 data to $L_{5100} < 10^{44.73} \text{ erg s}^{-1}$ and the S11 data to $L_{5100} > 10^{44.73} \text{ erg s}^{-1}$, in order to avoid overlap. The linear fit to the M_{BH} relation and a 1-1 relation are represented by solid and dotted lines respectively, while the zeropoint offset between the masses (Δ_{zpt}), intrinsic scatter with respect to a 1-1 relation (σ_{int} , dex), and the fraction where the masses overlap within error (p_{con}) are printed.

are consistent within uncertainty with those derived with and without the subtraction of the narrow $\text{FWHM}_{\text{MgII}}$ component. Therefore, we combined the literature data with respective treatment of the narrow Mg II and the averaged $\text{FWHM}_{\text{MgII}}$ in S11, to derive the $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{MgII}}$ relation altogether. We note that the relative FWHM offset of the literature data to our relation shown in Figure 13b are within the scatter of data points, indicating that the details of fitting to exclude the narrow Mg II component (S07; McGill et al. 2008; Dietrich et al. 2009) or to include but subtract it (Wang et al. 2009; S12) do not affect the line widths significantly when compared overall. The rms of all the data to our relation, 0.085 dex, is similar to that from Wang et al. (2009), and small enough to regard the $\text{FWHM}_{\text{MgII}}$ as a marginally good substitute of $\text{FWHM}_{\text{H}\beta}$ as much as $\text{FWHM}_{\text{H}\alpha}$.

Third, we derive the $\text{FWHM}_{\text{H}\beta}$ – FWHM_{CIV} relation in Figure 13c. The data can be fitted altogether with a log-linear relation, but the σ_{int} of the data to the relation, 0.212 dex, is large and comparable to the systematic uncertainty of single-epoch M_{BH} estimators when scaled as the FWHM squared. We checked the effect of fitting methodology by deriving all the L – L , FWHM–FWHM relations with the FITEXY method (Tremaine et al. 2002), to find that the slope of the $\text{FWHM}_{\text{H}\beta}$ – FWHM_{CIV} relation, 1.798 ± 0.026 , shows a meaningfully large difference to the BCES results. Still, we keep the BCES slope since it brings the Balmer and C IV M_{BH} estimates more consistent (section 4.3), and we note that the large σ_{int} between the $\text{FWHM}_{\text{H}\beta}$ – FWHM_{CIV} relation represents a poor correlation at best.

Overall, we find the relations of $\text{FWHM}_{\text{H}\alpha}$, $\text{FWHM}_{\text{MgII}}$ and FWHM_{CIV} against the $\text{FWHM}_{\text{H}\beta}$ without any noticeable evolution for the samples considered, covering a wide range of luminosity or redshift. Therefore, although the data for calibration is missing at $z \gtrsim 3$, we use our FWHM relations to calibrate the M_{BH} estimators for general usage in terms of redshift.

Our derived FWHM relations are as follows,¹⁷

$$\begin{aligned} \log \text{FWHM}_{\text{H}\beta,3} &= (1.061 \pm 0.013) \log \text{FWHM}_{\text{H}\alpha,3} + (0.055 \pm 0.008) \\ &= (1.226 \pm 0.032) \log \text{FWHM}_{\text{MgII},3} - (0.078 \pm 0.021) \\ &= (1.054 \pm 0.057) \log \text{FWHM}_{\text{CIV},3} - (0.024 \pm 0.045). \end{aligned} \quad (4)$$

We note the possibility that the FWHM relations established with spectra taken in different epochs could be affected by variability of the emission line shape, especially when the dynamic range of the probed FWHMs are narrower than that of the continuum luminosities. Still, the offset or scatter of the references with and without simultaneously acquired data in Figure 13 do not differ with each other. Also, we find the average and rms scatter in the ratio of the SDSS-III BOSS over SDSS-I/SDSS-II FWHM_{CIV} from 12, S/N > 15 objects matched with the *AKARI* sample to be 0.99 ± 0.08 , which implies the effect of variability to the shift or broadening of the $\text{FWHM}_{\text{H}\beta}$ – FWHM_{CIV} relation is negligible compared to the σ_{int} of the relation by an order of magnitude.

4.3. Updated M_{BH} estimators

Using the scaling relations obtained above, we now present an updated set of M_{BH} estimators based on various lines. Since the reverberation mapping of H β and L_{5100} forms the basis of mass estimation for AGNs, we start from the M_{BH} estimator that uses the 5100 Å luminosity and H β line width to take the following form and derive other estimators using scaling relations discussed above:

$$M_{\text{BH}} = \frac{f}{G} R_{\text{BLR}}(L_{5100}) \left(\frac{\text{FWHM}_{\text{H}\beta}}{2} \right)^2. \quad (5)$$

Previous secondary calibrations to the M_{BH} were performed through replacing the $(L_{5100}, \text{FWHM}_{\text{H}\beta})$ of the local reverberation mapped sample or SDSS AGNs, by $(L_{\text{H}\alpha}, \text{FWHM}_{\text{H}\alpha})$ (e.g., Greene & Ho 2005), $(L_{3000}, \text{FWHM}_{\text{MgII}})$ (e.g., McLure & Jarvis 2002), or $(L_{1350}, \text{FWHM}_{\text{CIV}})$ (e.g., Vestergaard & Peterson 2006). Using the latest R – L relation from Bentz et al. (2013)¹⁸

¹⁷ Throughout this paper we use subscript numbers to the line width to indicate its unit, such as $\text{FWHM}_{\text{H}\beta,3} = \text{FWHM}_{\text{H}\beta}/10^3 \text{ km s}^{-1}$.

¹⁸ We shifted the relation to Hubble parameter of $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ which is our adopted value throughout this work.

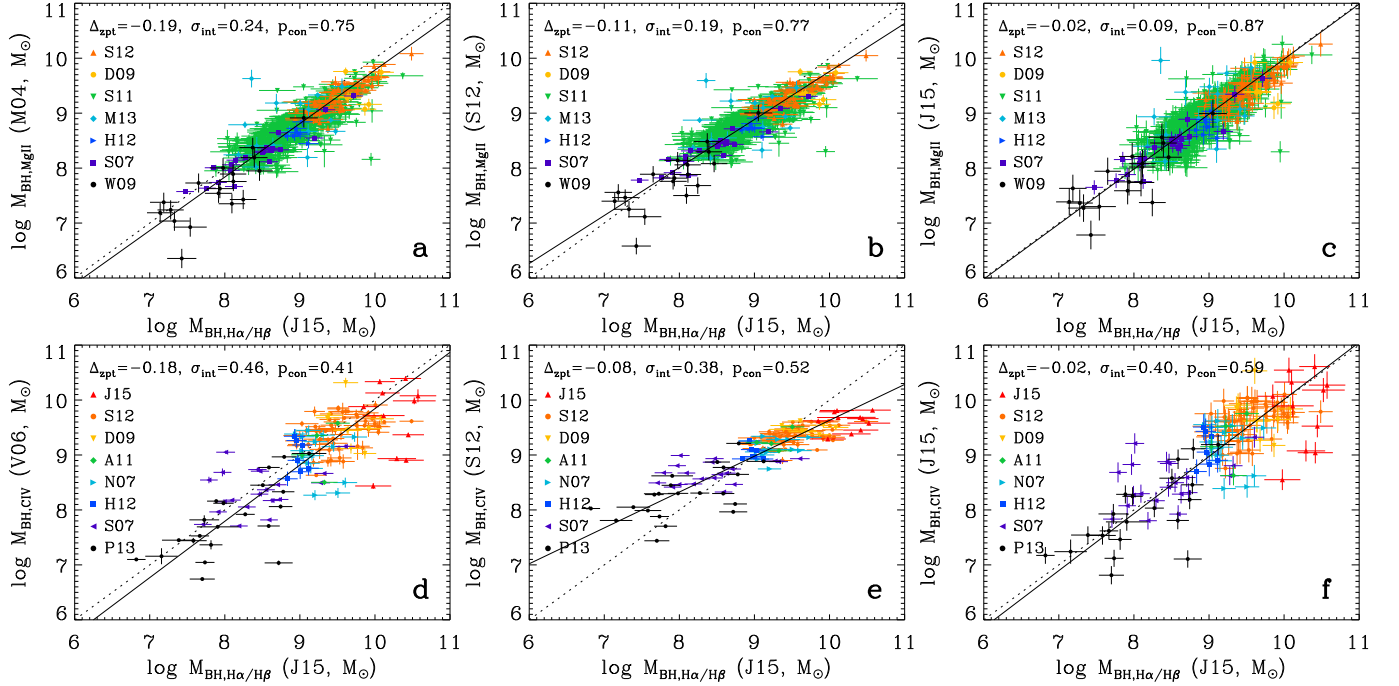


FIG. 15.— The comparison of Balmer M_{BH} to the UV M_{BH} of AGNs, from combined references. The references abbreviated on the plot are summarized in Table 4, with M04, V06 estimators from McLure & Dunlop (2004) and Vestergaard & Peterson (2006). All measurements are converted to our adopted cosmology and f -factor. We use the $M_{\text{BH}}(L_{5100}, \text{FWHM}_{\text{H}\alpha})$ for the Balmer masses, unless when the $\text{H}\alpha$ line was uncovered and the $M_{\text{BH}}(L_{5100}, \text{FWHM}_{\text{H}\beta})$ was used. The rest of the figure format follows that of Figure 14. We removed the objects unused in Figures 13 and 14.

and the constant for the mass equation (f -factor, $f = 5.1 \pm 1.3$) from Woo et al. (2013), we derive first the $\text{H}\beta$ M_{BH} estimator as,

$$R_{\text{BLR}} = (34.7 \pm 2.5) L_{5100,44}^{(0.533 \pm 0.034)} \text{lt} - \text{day}$$

$$M_{\text{BH}}(L_{5100}, \text{FWHM}_{\text{H}\beta}) = (8.63 \pm 2.29) \times 10^6 \quad (6)$$

$$\times L_{5100,44}^{(0.533 \pm 0.034)} \text{FWHM}_{\text{H}\beta,3}^2 M_{\odot}.$$

Replacing the $\text{H}\beta$ line width through the $\text{FWHM}_{\text{H}\beta}$ – $\text{FWHM}_{\text{H}\alpha}$ relation from Equation (4) we get the M_{BH} from $(L_{5100}, \text{FWHM}_{\text{H}\alpha})$,

$$M_{\text{BH}}(L_{5100}, \text{FWHM}_{\text{H}\alpha}) = (1.11 \pm 0.30) \times 10^7 \quad (7)$$

$$\times L_{5100,44}^{(0.533 \pm 0.034)} \text{FWHM}_{\text{H}\alpha,3}^{(2.12 \pm 0.03)} M_{\odot}.$$

Finally, the application of Equation (2) yields the M_{BH} from $(L_{\text{H}\alpha}, \text{FWHM}_{\text{H}\alpha})$,

$$M_{\text{BH}}(L_{\text{H}\alpha}, \text{FWHM}_{\text{H}\alpha}) = (5.20 \pm 1.41) \times 10^6 \quad (8)$$

$$\times L_{\text{H}\alpha,42}^{(0.511 \pm 0.033)} \text{FWHM}_{\text{H}\alpha,3}^{(2.12 \pm 0.03)} M_{\odot}.$$

Likewise, replacing the continuum luminosity and line width in Equation (6) to those of Mg II and C IV using Equations (3) and (4) yields,

$$M_{\text{BH}}(L_{3000}, \text{FWHM}_{\text{MgII}}) = (4.19 \pm 1.19) \times 10^6 \quad (9)$$

$$\times L_{3000,44}^{(0.548 \pm 0.035)} \text{FWHM}_{\text{MgII},3}^{(2.45 \pm 0.06)} M_{\odot}$$

$$M_{\text{BH}}(L_{1350}, \text{FWHM}_{\text{CIV}}) = (4.72 \pm 1.63) \times 10^6 \quad (10)$$

$$\times L_{1350,44}^{(0.547 \pm 0.037)} \text{FWHM}_{\text{CIV},3}^{(2.11 \pm 0.11)} M_{\odot}.$$

The M_{BH} 's based on $\text{H}\alpha$ and C IV from our estimators for the *AKARI* sample are given in Table 3. It

is worth noting that Equation (5) may not hold if the $\text{FWHM}_{\text{H}\beta}$ is not exactly proportional to the velocity dispersion, σ (Peterson et al. 2004; Collin et al. 2006), or if the R – L relation breaks down at high luminosity where the relation has not been tested extensively with the reverberation technique ($L_{5100} > 10^{46} \text{ergs s}^{-1}$, Bentz et al. 2013). Nevertheless, considering the advantages of using the FWHM (to be robust under poor sensitivity, wings in the line profile, or deblending, see S12) for the single epoch mass estimation, and the current expectations in the high luminosity R – L relation (see section 5.1), our set of calibrations have the merit where the rest-UV to optical M_{BH} estimations are mutually consistent through a wide range of redshifts, luminosities, and fitting methodologies.

To check the consistency between our M_{BH} estimators, we first compare the Balmer M_{BH} using the estimator in this work and existing estimators in Figure 14. The $M_{\text{BH,H}\beta}$ and $M_{\text{BH,H}\alpha}$ are compared, where a $\Delta M_{\text{BH}} < 0.3 \text{dex}$ error cut is applied to the data out of various references. The M_{BH} values are derived from the G05, S11 estimators and from this work (Equations (6), (7)) in Figures 14a–14c respectively, where the zero-point offset between the $\text{H}\alpha$ and $\text{H}\beta$ masses (Δ_{zpt}), intrinsic scatter with respect to a 1-1 relation (σ_{int}), and the fraction of the data points where 1σ errors of M_{BH} values from different estimators overlap with each other (p_{con}) are shown. We find that the masses from our estimators are closer to a 1-1 relation than of G05 or S11, throughout the range 10^6 – $10^{10} M_{\odot}$. Also, fully considering the propagated errors in the mass equation, our estimators may overestimate the M_{BH} uncertainty when comparing the $\text{H}\beta$ and $\text{H}\alpha$ masses as reflected from the negative σ_{int} , which suggests that the actual difference in the

TABLE 5
 L_{5100} – $L_{H\alpha}$ RELATION

Reference (1)	α (2)	β (3)	σ_{int} (4)	$\sigma_{\text{int,all}}$ (5)	N (6)	z (7)	$\log L_{5100}$ (8)	Dataset (9)
Greene & Ho (2005)	0.720 ± 0.002	1.157 ± 0.005	0.078	0.241	229	< 0.35	41.7–45.0	G05
Shen & Liu (2012)	0.791 ± 0.093	1.010 ± 0.042	0.088	0.153	60	1.5–2.2	45.4–47.0	S12
This work	0.646 ± 0.011	1.044 ± 0.008	0.095	0.095	464	0.0–6.2	41.7–47.2	G05, S07, S11, H12, S12, M13, J15

NOTE. — The L_{5100} – $L_{H\alpha}$ relations from previous studies and this work, where α and β are defined as $\log L_{H\alpha,42} = \alpha + \beta \log L_{5100,44}$. Column 1: Reference; Column 2: α and its uncertainty (1σ); Column 3: β and its uncertainty; Column 4: Intrinsic scatter (dex) of the relation over the L_{5100} range covered; Column 5: Intrinsic scatter of the relation over the range $41.7 < \log L_{5100} < 47.2$; Column 6: Number of objects used; Column 7: Redshift range; Column 8: Range of L_{5100} (ergs s $^{-1}$); Column 9: References of AGN dataset with abbreviations from Table 4.

TABLE 6
VIRIAL MASS ESTIMATORS

Reference (1)	α (2)	β (3)	γ (4)	N (5)	z (6)	$\log M_{\text{BH}}$ (7)	Method (8)
$M_{\text{BH}}(L_{5100,44}, \text{FWHM}_{H\beta,3})$							
Greene & Ho (2005)	6.64 ± 0.02	0.64 ± 0.02	2	35	< 0.37	5.5–9.0	RM
Vestergaard & Peterson (2006)	6.91 ± 0.02	0.5	2	25	0.00–0.29	7.2–9.3	RM
Bentz et al. (2013), this work	6.94 ± 0.12	0.533 ± 0.034	2	41	0.00–0.29	6.0–10.7	RM
$M_{\text{BH}}(L_{5100,44}, \text{FWHM}_{H\alpha,3})$							
Greene & Ho (2005)	6.70 ± 0.06	0.64 ± 0.02	2.06 ± 0.06	162	< 0.37	5.2–9.0	SE
Shen & Liu (2012)	7.01	0.555	1.87	60	1.5–2.2	8.8–10.3	SE
This work	7.05 ± 0.12	0.533 ± 0.034	2.12 ± 0.03	654	0.0–2.4	5.8–10.6	SE
$M_{\text{BH}}(L_{H\alpha,42}, \text{FWHM}_{H\alpha,3})$							
Greene & Ho (2005)	6.30 ± 0.08	0.55 ± 0.02	2.06 ± 0.06	243	< 0.37	5.1–9.0	SE
Shen & Liu (2012)	6.55	0.564	1.82	60	1.5–2.2	8.8–10.2	SE
This work	6.72 ± 0.12	0.511 ± 0.033	2.12 ± 0.03	969	0.0–6.2	5.6–10.5	SE
$M_{\text{BH}}(L_{3000,44}, \text{FWHM}_{\text{MgII},3})$							
McLure & Dunlop (2004)	6.5	0.62	2	22	0.03–0.37	7.5–8.9	RM
Vestergaard & Osmer (2009)	6.86	0.5	2	—	—	—	SE
Wang et al. (2009)	7.15 ± 0.27	0.46 ± 0.08	1.48 ± 0.49	29	0.00–0.16	6.3–9.0	RM
Shen & Liu (2012)	6.95	0.584	1.71	60	1.5–2.2	8.8–10.1	SE
This work	6.62 ± 0.12	0.548 ± 0.035	2.45 ± 0.06	1010	0.0–6.4	6.8–10.3	SE
$M_{\text{BH}}(L_{1350,44}, \text{FWHM}_{\text{CIV},3})$							
Vestergaard & Peterson (2006)	6.66 ± 0.01	0.53	2	27	0.00–0.23	5.5–9.3	RM
Shen & Liu (2012)	8.02	0.471	0.24	60	1.5–2.2	8.9–9.6	SE
Park et al. (2013)	7.48 ± 0.24	0.52 ± 0.09	0.56 ± 0.48	25	0.01–0.23	7.0–9.0	RM
This work	6.67 ± 0.15	0.547 ± 0.037	2.11 ± 0.11	258	0.0–5.4	6.8–10.7	SE

NOTE. — The M_{BH} relations from previous studies and this work, where α , β , γ are defined as $\log M_{\text{BH}} = \alpha + \beta \log L + \gamma \log \text{FWHM}$. Column 1: Reference; Column 2: α and its uncertainty (1σ); Column 3: β and its uncertainty; Column 4: γ and its uncertainty; Column 5: Number of objects used for luminosity or line width calibration; Column 6: Redshift range of objects used for luminosity or line width calibration; Column 7: Range of M_{BH} from reference; Column 8: Method of calibration, where RM, SE denote calibrations of the luminosity and line width using the reverberation mapped and single epoch samples, respectively.

f -factor or R_{BLR} between the $H\beta$ and $H\alpha$ line emitting regions are likely to be smaller than their uncertainty. Therefore, we regard the $H\beta$ and $H\alpha$ M_{BH} from our estimators to be indistinguishable, which is supported by the high fraction (98%) of $H\beta/H\alpha$ masses to be consistent within measurement error.

With the $H\beta$ and $H\alpha$ M_{BH} recipes checked to be mutually consistent, we further compare the Balmer to the UV-based M_{BH} with $\Delta M_{\text{BH}} < 0.3$ dex, in Figure 15. We use $M_{\text{BH},H\alpha}$ as the Balmer mass due to the stronger $H\alpha$ emission than $H\beta$ and the availability of $H\alpha$ line in our *AKARI* data, unless only the $M_{\text{BH},H\beta}$ can be estimated using the measurements in the literature. Compared to the conventional estimators calibrated at relatively low luminosity and redshift (McLure & Dunlop 2004; Vestergaard & Peterson 2006) or relatively high lu-

minosity and redshift (S12), our calibrations (Equations (6), (7), (9), (10)) bring the rest-UV (Mg II in Figures 15a–15c, C IV in Figures 15d–15f) and Balmer M_{BH} 's to be closer to a 1–1 relation in the range 10^7 – $10^{10} M_{\odot}$ in terms of both zeropoint offset and intrinsic scatter. Also, we find that our M_{BH} estimators fully considering the propagated errors give more realistic value of the error bars compared to estimators that do not, and raise the p_{con} . Moreover, previous UV-based M_{BH} 's are tilted to ours largely from the different slope of the FWHM term to that of our estimator (e.g., 0.24 in S12, 2.11 in this work for C IV, see also, Park et al. 2013), and this produces a systematic offset in M_{BH} values. For example, we find that S12 estimators produce a systematic difference between C IV and Balmer M_{BH} in such a way that $M_{\text{BH},\text{CIV}}$ are larger by 0.7 dex at $M_{\text{BH},H\alpha/H\beta} = 10^7 M_{\odot}$

and smaller by 0.4 dex at $M_{\text{BH,H}\alpha/\text{H}\beta} = 10^{10} M_{\odot}$ (Figure 15e). Previous UV-based M_{BH} calibrations from limited dynamic range yields σ_{int} to the Balmer M_{BH} as small as our estimator, and may place the C IV and Balmer masses to be consistent within a narrow range of masses, but we caution on the usage of conventional estimators when extensively comparing the rest-optical and UV M_{BH} .

Even when using our rest-UV mass equations with minimized zeropoint and systematic offsets from a large dynamic range, the intrinsic scatter is another issue. Although our σ_{int} of the $M_{\text{BH,MgII}}$ to the Balmer M_{BH} is relatively small (0.09 dex), the σ_{int} of the $M_{\text{BH,CIV}}$ to the Balmer M_{BH} (0.40 dex) is comparable to the systematic uncertainty of the single epoch mass estimator itself (Bentz et al. 2013; Woo et al. 2013). Moreover, since our estimators take into account the measurement error of f -factor, $R-L$ relation, scatter in the UV-optical luminosity and line width relations, the non-negligible fraction of rest-UV masses that deviate from the Balmer masses more than their overlapping measurement errors (13% of the $M_{\text{BH,MgII}}$ and 41% of the $M_{\text{BH,CIV}}$) suggest that one needs to be cautious about the usage of the C IV line when deriving M_{BH} values.

5. DISCUSSION

5.1. Reliability of single-epoch AGN black hole masses

Does the relations between continuum and line luminosities, and the FWHM relations that hold up to the highest luminosity AGNs, guarantee the accurate mass estimation of the most luminous AGNs? Unfortunately the answer is not clear yet, since the $R_{\text{BLR,H}\beta}-L_{5100}$ relation is not observationally probed for $L_{5100} > 10^{46} \text{ erg s}^{-1}$ (Bentz et al. 2013). A hint to estimate the high luminosity end behavior of the optical $R-L$ relation is to look at the $R_{\text{BLR,CIV}}-L_{1350}$ relation, where the $R_{\text{BLR,CIV}}$ traces the inner part to the H β line region. Although the current number of C IV reverberation measurements is small (Kaspi et al. 2007, Sluse et al. 2011, Chelouche et al. 2012), they cover up to luminosities of $L_{1350} = 10^{47.0} \text{ erg s}^{-1}$ and $L_{5100} = 10^{46.8} \text{ erg s}^{-1}$. The slope of the relation in Kaspi et al. (2007), $(0.52-0.55) \pm 0.04$, is within the uncertainty to the slope of the $R_{\text{BLR,H}\beta}-L_{5100}$ relation (Bentz et al. 2013), 0.533 ± 0.034 . This suggests that the ratio of $R_{\text{BLR,H}\beta}$ -to- $R_{\text{BLR,CIV}}$ is nearly a constant at 2-4 over $L_{5100} < 10^{46} \text{ erg s}^{-1}$, where the proportionality of L_{5100} and L_{1350} suggests that it is likely so at a higher luminosity range. Therefore, we do not expect a break in the luminous end optical $R-L$ relation unless the optical data will somehow fail to form a simple relation at high luminosity where the current UV relation holds. Future compilations of both the C IV and H β broad line region sizes will help to better constrain the luminous end $R-L$ relation.

Even if we find it plausible to assume that the $R_{\text{BLR,H}\beta}-L_{5100}$ relation does not change its slope in the luminous end, the large scatter in between the Balmer and rest-UV M_{BH} imposes further limitations on the usage of UV M_{BH} estimators. As discussed in the introduction, it has been known that there is a large σ_{int} in the C IV and Balmer M_{BH} relation in comparison to that of the Mg II and Balmer. Our comparison of $M_{\text{BH,H}\alpha}$

against $M_{\text{BH,CIV}}$ for high luminosity AGNs at $z > 3.3$ further shows consistent results. Indeed, we find 4 out of 11 $M_{\text{BH,CIV}}$ of *AKARI* quasars in Figure 15f to be scattered to the $M_{\text{BH,H}\alpha}$ by an order of magnitude or larger, therefore the use of C IV masses can be questioned up to the highest redshifts even when considering the additional uncertainties in the *AKARI* $M_{\text{BH,H}\alpha}$ tested in Figure 5. Overall, our result supports that the M_{BH} estimate from C IV contains a rather large uncertainty at any redshift or luminosity.

The greatest uncertainties when calibrating the UV M_{BH} estimators in this work ($\sigma_{\text{int}} > 0.1$ dex), come from the scatter in the $L_{5100}-L_{1350}$ and $\text{FWHM}_{\text{H}\beta}-\text{FWHM}_{\text{CIV}}$ relations. Since a variety of obscuration in the rest-UV continuum of quasars may result in a scattered $L_{5100}-L_{1350}$ relation, we suggest to take into account the level of obscuration when establishing the $M_{\text{BH,CIV}}$ estimator, although careful treatment is required as it is controversial whether the color correction into the M_{BH} works to reduce the scatter between the C IV and Balmer masses (Assef et al. 2011; S12). Meanwhile, the scatter in the FWHM relations which is the largest for the $\text{FWHM}_{\text{H}\beta}-\text{FWHM}_{\text{CIV}}$ relation, could arise from broad absorption features or a non-reverberating component of the C IV broad emission (Denney 2012), where future reverberation mapping will further enable us to test. Without detailed analysis on the origin of the scatter in between the Balmer and rest-UV M_{BH} , we caution on the individual measurement of M_{BH} from UV, especially the one using C IV.

5.2. On the massive end black hole mass evolution

Having cross-calibrated the M_{BH} estimators, we plot the Balmer and Mg II M_{BH} measurements of quasars along redshift in Figure 16a, from our *AKARI* observations and the literature indicated on the figure. First of all, we notice that the most massive envelope that stays at $\sim 10^{10} M_{\odot}$ up to $z \sim 5$, starts to disappear at higher redshift. To quantify how significant the M_{BH} 's are evolving at the massive end for $z > 5$ quasars, we performed the Kolmogorov-Smirnov (K-S) test to compare the mass distribution at above and below $z = 5$, at $M_{\text{BH}} > 2 \times 10^9 M_{\odot}$. This mass threshold for completeness is governed by the shallow *AKARI* data, through the $L_{5100} \gtrsim 10^{46.5} \text{ erg s}^{-1}$ limit at $\text{S/N}_{\text{H}\alpha} = 3$ and the FWHM limit of 2500 km s^{-1} from section 3.1. The K-S probability $p_{\text{KS}} = 0.34\%$ turns out to be meaningfully small, supporting the massive end BHs at $z < 5$ to be heavier than at $z > 5$. We checked our results to be unchanged much by the inclusion of *AKARI* data that could bias the masses by up to ~ 0.1 dex (section 3.1), by performing the same test while giving a -0.1 dex correction to the *AKARI* data, only to find the trend to be mildly weaker ($p_{\text{KS}} = 0.64\%$).

We cross-check if this evolution in M_{BH} is also reflected through the M_{BH} histograms. In order to correct for the luminosity selection effect in number counting, we plot the 1450 \AA absolute magnitude (M_{1450}) histogram in Figure 16b. We overplot the quasar luminosity functions in redshift bins of $2 < z < 3$, $3 < z < 5$, $z > 5$, scaled by a constant so that they match the M_{1450} counts in total number at $M_{1450} < -26.4$. We use only the $M_{1450} < -26.4$ data for better completeness in scal-

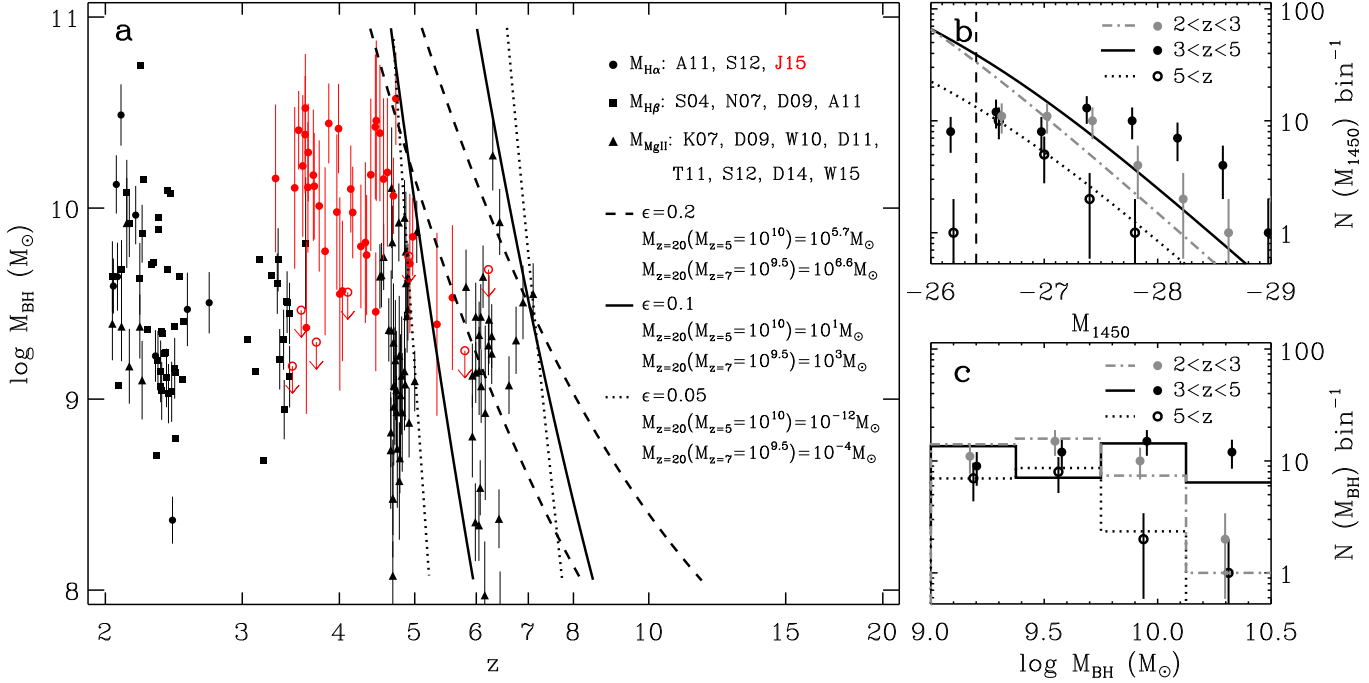


FIG. 16.— The M_{BH} of AGNs along redshift (left). The M_{BH} from $\text{H}\alpha$ are in circles, $\text{H}\beta$ in squares, and Mg II in triangles, while our *AKARI* data points are in red circles. For objects with multiple lines, we choose the $\text{H}\alpha$ over $\text{H}\beta$, and $\text{H}\beta$ over Mg II , while for repeated measurements of an object we used the latest results. Most of the references abbreviated on the plot are summarized in Table 4, with S04, K07, W10, D11, T11, D14, W15 additionally from Shemmer et al. (2004), Kurk et al. (2007), Willott et al. (2010a), De Rosa et al. (2011), Trakhtenbrot et al. (2011), De Rosa et al. (2014), and Wu et al. (2015). We do not plot the error for the S04 data as they are not available, and do not use a part of the T11 data with poor quality (flagged “3”). The model tracks of exponential BH growth that matches the massive envelope of AGNs at $z \sim 5$ and 7 are plotted, with each set of parameters describing the curve shown. On the right, we plot the histogram of M_{1450} (top) and M_{BH} (bottom) at each denoted redshift bin in circles, with errors determined through Poisson statistics. The luminosity functions scaled to the observed number counts to yield the same number at $M_{1450} < -26.4$ at each redshift (top), and the resultant mass counts from the luminosity scaling and a $\text{FWHM} > 2500 \text{ km s}^{-1}$ cut (bottom), are overplotted in lines with styles marked next to the redshift.

ing the luminosities. The break luminosity (M_{1450}), the faint end slope (α), and the bright end slope (β) of the luminosity function are fixed to $M_{1450} = -26.39$ at $z > 3$ and $M_{1450} = -25.5$ at $z < 3$ while $(\alpha, \beta) = (-1.80, -3.26)$ at all redshifts (McGreer et al. 2013; Ross et al. 2013). The set of (α, β) are fixed as the evolution in the α value does not affect the luminosity regime of interest much, and β stays nearly constant with redshift. Overall, we find that the quasar samples at $z < 5$ are in excess of luminous objects compared to the luminosity functions.

Not only there is a luminosity selection effect in the quasars plotted in Figure 16a, there is a FWHM limit for each reference data which is the poorest for the *AKARI* data. Taking into account of these effects, we plot the M_{BH} histograms in Figure 16c before (points) and after (histograms) applying the luminosity selection correction and a $\text{FWHM} = 2500 \text{ km s}^{-1}$ cut. After correcting the luminosity selection and giving the FWHM cut the M_{BH} distribution is shifted to more $10^9 M_{\odot}$ BHs and less $10^{10} M_{\odot}$ BHs at $z < 5$. Still, the number of $10^{10} M_{\odot}$ BHs start to drop at $z > 5$, whereas the $10^9 M_{\odot}$ BHs do not exhibit an evolution up to Poisson error. Thus, we interpret the M_{BH} histogram that there is an evolution in the number of $10^{10} M_{\odot}$ BHs to be increasing with time at $z > 5$. The $z > 5$, $\sim 10^9 M_{\odot}$ AGNs are likely to be in a rapidly growing state such that their masses can soon reach $\sim 10^{10} M_{\odot}$, consistent with the higher Eddington ratio trend found in distant, luminous AGNs (Willott et al. 2010a; De Rosa et al. 2011). Note that the number density of quasars evolve strongly from $z = 2$

to 6 (e.g., Willott et al. 2010a). We caution that the mass and luminosity histograms in Figures 16b and 16c do not reflect this number density evolution, and that Figures 16b and 16c are used for comparing the shape of these functions. The fact that the numbers of $z = 2$ through 6 quasars appear similar in these figures is due to the fact that the sampling of quasars is different at each redshift bin.

We further investigate the upper mass envelope of $z > 5$ AGNs to check if their extremely high masses can be explained by the earliest BH growth from model seed masses. For this we follow Volonteri & Rees (2005) to assume a continuous, Eddington limited accretion of matter to a BH seed at $z = 20$, where the BH will grow with time as Equation (1), with $t_0 = t_{z=20}$ and $M_0 = M_{z=20}$ for the age of the universe and the seed BH mass, respectively at $z = 20$. For each given $\epsilon = 0.05, 0.1, 0.2$, we determined a pair of seed masses M_0 such that the growth curve reaches the observed massive limit at $z \sim 5$ and $z \sim 7$ respectively, shown in Figure 16a.

When $\epsilon = 0.05$, the seed masses to explain the most massive BHs at $z \sim 5$ and 7 are $M_0 \sim 10^{-12} M_{\odot}$ and $10^{-4} M_{\odot}$ respectively, which in other words, gives plenty of time for seed BHs of any meaningful mass to grow up to the most massive quasars without a strictly continuous, Eddington limited accretion. One can expect to find fully grown massive BHs at $z \sim 6$ detectable as high or low Eddington ratio AGNs, if such rapid growth is possible. Luminous AGNs at $z \sim 6$ have high Eddington ratios so far, but future studies may uncover fainter objects

with weakly active BHs. If $\epsilon = 0.1$ that is roughly consistent with the Soltan argument measurements (Soltan 1982; Yu & Tremaine 2002), it predicts $M_0 \sim 10^{1-3} M_\odot$ which gives a reasonable estimate of the seed mass to be Population III stars. However, it may be difficult for stellar seed BHs to keep its Eddington limited accretion from $z = 20$ to $7-5$, which is $0.6-1.0$ Gyrs in duration and is longer than a typical quasar lifetime ($\lesssim 0.1$ Gyr, Hopkins et al. 2005), perhaps requiring extended period of accretion or BH-BH mergers. Lastly, the $\epsilon = 0.2$ model which corresponds to the case when BHs are rotating rapidly (e.g., Kerr 1963; Marconi et al. 2004) only accepts very heavy seed BHs at $M_0 \sim 10^{5.7-6.6} M_\odot$, requiring supercritical accretion from lighter seed masses (Volonteri & Rees 2005) or supports direct collapse of massive primordial gas (Volonteri et al. 2008). With $\epsilon = 0.2$, we expect that the most massive $\sim 10^{10} M_\odot$ BHs from $M_0 \sim 10^6 M_\odot$ seeds can only start appearing at the redshifts between 5 and 6.

6. SUMMARY

We measured the redshifted $H\alpha$ emission and rest-UV to rest-optical continuum properties of 155 luminous quasars at $3 \lesssim z \lesssim 6$ using the *AKARI* spectra and other datasets, and estimated their M_{BH} . We summarize our findings as the following.

1. The $L_{5100}-L_{H\alpha}$ relation holds up to the most luminous quasars ($L_{5100} \sim 10^{47} \text{ erg s}^{-1}$) with a single slope unchanging up to $z \sim 6$, suggesting a consistent response of the broad line region to the incident continuum irrespective of AGN luminosity and redshift.
2. The relations between rest-optical and UV continuum and line luminosities, and the FWHM relations hold up to the highest luminosity AGNs. Together with predictions of an extended optical $R-L$ relation to the highest luminosities, it enables us to calibrate the rest-UV M_{BH} estimators to be consistent with the Balmer masses overall. However, some of the rest-UV and optical M_{BH} are scattered more than their uncertainties including the errors from the recipe, for only 13% of Mg II masses but for 41% of the C IV estimation. The C IV masses have a 0.40 dex intrinsic scatter to the Balmer masses, which places negative implications on its reliability.
3. The massive end envelope of M_{BH} steeply evolves at $z > 5$, suggesting they are in a rapidly growing state from given seed masses. The most massive BHs at $z = 5-7$ can be explained by the Eddington limited, continuous accretion onto $\sim 10^{1-3} M_\odot$ seed masses at $z = 20$ if $\epsilon = 0.1$ and the formation redshift is $z = 20$ for the seed BH, while there are a range of viable accretion rates and seed masses if ϵ is different.

We expect future observations to compile sensitive measurements of ($L_{5100}, L_{H\alpha}$) to better identify the outliers in the $L_{5100}-L_{H\alpha}$ relation, where an example would be a small population of weak emission line quasars. Also, future rest-optical reverberation mapping of high luminosity AGNs will verify if the prediction on the $R-L$ relation to extend with a single slope will hold. Most importantly, the discovery of the highest redshift quasars will further uncover the evolutionary tracks of the earliest BH growth, also improving the current understanding of M_{BH} growth at $z > 5$ from small number statistics.

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TABLE 7
CONTINUUM AND LINE BASED PROPERTIES OF THE SAMPLE

Name (1)	z_{ref} (2)	$z_{\text{H}\alpha}$ (3)	$\log L_{1350}$ (4)	$\log L_{5100}$ (5)	$\log L_{\text{H}\alpha}$ (6)	$\text{FWHM}_{3, \text{CIV}}$ (7)	$\text{FWHM}_{3, \text{H}\alpha}$ (8)	$\log M_{\text{BH}, \text{CIV}}$ (9)	$\log M_{\text{BH}, \text{H}\alpha}$ (10)
SDSS J000239.39+255034.8	5.80	5.79	99.00 ± 99.00	46.56 ± 0.02	45.14 ± 0.11	99.00 ± 99.00	2.50 ± -1.00	99.00 ± 99.00	9.25 ± -1.00
Q 0000-26	4.10	4.11	47.40 ± 0.01	47.13 ± 0.01	45.48 ± 0.13	4.33 ± 0.28	2.50 ± -1.00	9.88 ± 0.22	9.56 ± -1.00
SDSS J000552.34-000655.8	5.85	5.85	99.00 ± 99.00	46.04 ± 0.08	44.75 ± 0.32	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR J0006-6208	4.45	4.49	46.90 ± 0.02	46.71 ± 0.00	45.46 ± 0.06	11.33 ± 1.33	2.85 ± 0.85	10.48 ± 0.24	9.46 ± 0.31
SDSS J001115.23+144601.8	4.97	99.00	99.00 ± 99.00	46.78 ± 0.02	47.97 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR J0018-3527	4.15	4.15	99.00 ± 99.00	46.72 ± 0.01	45.39 ± 0.29	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
PMN J0022-0759	3.90	99.00	99.00 ± 99.00	46.57 ± 0.01	47.97 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR 0019-1522	4.53	99.00	46.95 ± 0.02	46.58 ± 0.02	47.92 ± -1.00	3.48 ± 0.33	99.00 ± 99.00	9.43 ± 0.21	99.00 ± 99.00
BR J0030-5129	4.17	99.00	99.00 ± 99.00	46.55 ± 0.01	48.02 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BRI J0048-2442	4.15	99.00	99.00 ± 99.00	99.00 ± 99.00	48.08 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BRI J0048-244	4.15	99.00	99.00 ± 99.00	46.07 ± 0.03	47.90 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J010619.24+004823.3	4.45	4.44	46.99 ± 0.01	46.80 ± 0.02	45.69 ± 0.19	3.00 ± 0.20	7.75 ± 2.92	9.32 ± 0.11	10.43 ± 0.38
BRI J0113-280	4.30	99.00	99.00 ± 99.00	46.38 ± 0.01	47.92 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J011351.96-039551.1	3.67	3.68	46.93 ± 0.03	46.54 ± 0.01	47.65 ± -1.00	8.64 ± 0.36	99.00 ± 99.00	10.25 ± 0.22	99.00 ± 99.00
SDSS J011453.81-103934.0	3.73	99.00	46.65 ± 0.12	46.62 ± 0.01	47.69 ± -1.00	5.11 ± 0.35	99.00 ± 99.00	9.62 ± 0.22	99.00 ± 99.00
SDSS J012211.11+150914.3	4.46	99.00	47.00 ± 0.04	46.58 ± 0.02	47.91 ± -1.00	7.59 ± 0.82	99.00 ± 99.00	10.17 ± 0.23	99.00 ± 99.00
SDSS J012403.77+004432.7	3.83	3.83	47.11 ± 0.01	46.83 ± 0.01	45.44 ± 0.20	8.06 ± 0.29	3.76 ± 1.69	10.29 ± 0.15	9.77 ± 0.44
SDSS J012700.68-004559.1	4.08	4.07	47.01 ± 0.01	46.68 ± 0.01	45.14 ± 0.33	6.96 ± 0.27	99.00 ± 99.00	10.10 ± 0.13	99.00 ± 99.00
BRI J0137-422	3.97	3.98	99.00 ± 99.00	46.42 ± 0.01	45.49 ± 0.22	99.00 ± 99.00	5.95 ± 2.46	99.00 ± 99.00	9.98 ± 0.41
SDSS J014049.18-083942.5	3.71	3.69	47.24 ± 0.01	46.96 ± 0.01	45.51 ± 0.12	4.84 ± 0.07	5.05 ± 1.27	9.89 ± 0.21	10.11 ± 0.28
SDSS J015048.83+004126.2	3.70	3.70	46.95 ± 0.01	46.64 ± 0.01	45.42 ± 0.17	4.14 ± 0.34	6.47 ± 2.18	9.58 ± 0.11	10.17 ± 0.35
SDSS J021646.94-092107.3	3.72	99.00	47.01 ± 0.01	46.74 ± 0.01	45.47 ± 0.11	9.87 ± 0.96	99.00 ± 99.00	10.42 ± 0.18	99.00 ± 99.00
SDSS J023137.65-072854.4	5.41	99.00	46.66 ± 0.07	46.59 ± 0.02	48.20 ± -1.00	3.03 ± 0.33	99.00 ± 99.00	9.15 ± 0.22	99.00 ± 99.00
BR J0234-1806	4.31	99.00	99.00 ± 99.00	46.54 ± 0.01	47.99 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J024447.79-081606.0	4.07	99.00	47.19 ± 0.02	46.94 ± 0.01	47.96 ± -1.00	9.76 ± 0.54	99.00 ± 99.00	10.51 ± 0.23	99.00 ± 99.00
BR J0301-5537	4.13	99.00	99.00 ± 99.00	46.41 ± 0.01	47.86 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR J0307-4945	4.73	4.78	47.47 ± 0.01	47.12 ± 0.01	45.79 ± 0.11	6.40 ± 0.86	7.59 ± 1.51	10.27 ± 0.25	10.57 ± 0.24
BR J0324-2918	4.62	4.60	99.00 ± 99.00	46.94 ± 0.01	45.29 ± 0.27	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J034402.85-065300.6	3.96	99.00	46.98 ± 0.02	46.44 ± 0.01	47.83 ± -1.00	7.23 ± 0.24	99.00 ± 99.00	10.11 ± 0.21	99.00 ± 99.00
BR J0426-2202	4.32	4.33	47.13 ± 0.01	46.72 ± 0.01	45.46 ± 0.20	5.55 ± 1.16	4.21 ± 1.80	9.96 ± 0.28	9.82 ± 0.42
BR J0525-3343	4.38	99.00	47.01 ± 0.02	46.84 ± 0.01	47.89 ± -1.00	11.23 ± 2.14	99.00 ± 99.00	10.54 ± 0.28	99.00 ± 99.00
BR J0529-3526	4.41	4.41	99.00 ± 99.00	46.64 ± 0.01	45.35 ± 0.33	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BR J0714-6455	4.46	4.44	47.32 ± 0.01	46.84 ± 0.01	45.45 ± 0.23	11.42 ± 0.33	7.85 ± 3.33	10.72 ± 0.23	10.46 ± 0.42
SDSS J073149.51+285448.7	3.68	3.69	46.96 ± 0.02	46.50 ± 0.01	45.13 ± 0.28	7.08 ± 0.63	99.00 ± 99.00	10.09 ± 0.22	99.00 ± 99.00
SDSS J075303.33+423130.8	3.59	3.56	47.23 ± 0.01	46.79 ± 0.01	45.32 ± 0.19	2.93 ± 0.08	6.24 ± 2.29	9.42 ± 0.10	10.22 ± 0.37
SDSS J080430.57+542041.1	3.76	3.75	46.99 ± 0.01	46.74 ± 0.01	45.29 ± 0.24	7.13 ± 0.32	99.00 ± 99.00	10.11 ± 0.15	99.00 ± 99.00
SDSS J080849.43+521515.3	4.46	99.00	47.16 ± 0.01	46.81 ± 0.01	47.91 ± -1.00	8.41 ± 0.31	99.00 ± 99.00	10.35 ± 0.13	99.00 ± 99.00
SDSS J081754.52+413225.3	3.54	99.00	46.59 ± 0.02	46.40 ± 0.01	47.67 ± -1.00	4.94 ± 0.12	99.00 ± 99.00	9.55 ± 0.11	99.00 ± 99.00
SDSS J081827.40+172251.8	6.00	99.00	99.00 ± 99.00	46.74 ± 0.12	48.40 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J083118.52+424728.8	3.32	99.00	46.58 ± 0.04	46.27 ± 0.01	47.58 ± -1.00	7.21 ± 0.48	99.00 ± 99.00	9.90 ± 0.16	99.00 ± 99.00
SDSS J083700.82+350550.2	3.31	3.30	46.87 ± 0.02	46.62 ± 0.01	45.36 ± 0.19	4.03 ± 0.17	6.42 ± 2.50	9.52 ± 0.12	10.16 ± 0.39
SDSS J083946.22+511202.8	4.39	4.39	46.84 ± 0.03	46.71 ± 0.01	45.61 ± 0.21	5.78 ± 0.21	6.22 ± 2.50	9.83 ± 0.15	10.17 ± 0.40
SDSS J084119.52+290504.4	5.96	99.00	99.00 ± 99.00	46.18 ± 0.07	48.35 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J085837.95+052141.7	3.53	99.00	46.84 ± 0.03	46.53 ± 0.01	47.62 ± -1.00	6.82 ± 0.68	99.00 ± 99.00	9.98 ± 0.22	99.00 ± 99.00
87GB 090153.2+694215.5	5.47	99.00	99.00 ± 99.00	46.26 ± 0.03	48.11 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J090634.84+023433.8	4.51	99.00	47.03 ± 0.02	46.73 ± 0.02	47.99 ± -1.00	8.95 ± 1.18	99.00 ± 99.00	10.34 ± 0.19	99.00 ± 99.00
SDSS J092721.82+200123.7	5.77	99.00	99.00 ± 99.00	46.20 ± 0.06	48.32 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J093554.45+525616.5	4.01	3.99	46.90 ± 0.02	46.78 ± 0.01	45.43 ± 0.22	5.74 ± 0.37	3.04 ± 1.59	9.86 ± 0.21	9.55 ± 0.51
BR 0951-0450	4.37	99.00	46.71 ± 0.03	46.46 ± 0.02	48.03 ± -1.00	5.64 ± 0.60	99.00 ± 99.00	9.74 ± 0.22	99.00 ± 99.00
SDSS J095511.32+594030.7	4.34	4.30	47.03 ± 0.02	46.81 ± 0.01	45.72 ± 0.14	4.61 ± 0.26	3.71 ± 1.11	9.73 ± 0.15	9.75 ± 0.32
SDSS J095744.46+330820.7	4.23	99.00	47.17 ± 0.02	46.67 ± 0.01	48.05 ± -1.00	9.19 ± 0.67	99.00 ± 99.00	10.44 ± 0.23	99.00 ± 99.00
QUEST J101046.3-013104.1	5.09	99.00	99.00 ± 99.00	47.70 ± 0.01	48.02 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J101336.37+561536.4	3.63	3.64	46.95 ± 0.03	46.99 ± 0.09	45.58 ± 0.14	6.90 ± 0.91	6.65 ± 1.88	10.06 ± 0.24	10.38 ± 0.31
SDSS J102622.89+471907.0	4.94	99.00	99.00 ± 99.00	46.64 ± 0.01	47.97 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J103221.11+092749.0	4.00	4.02	47.02 ± 0.01	46.55 ± 0.02	45.31 ± 0.33	7.58 ± 0.43	99.00 ± 99.00	10.18 ± 0.13	99.00 ± 99.00
SDSS J103242.71+433605.4	3.46	99.00	46.50 ± 0.07	46.40 ± 0.01	47.62 ± -1.00	5.97 ± 0.19	99.00 ± 99.00	9.68 ± 0.15	99.00 ± 99.00
SDSS J104351.20+650647.6	4.47	99.00	46.85 ± 0.05	46.46 ± 0.01	47.91 ± -1.00	4.07 ± 1.14	99.00 ± 99.00	9.52 ± 0.32	99.00 ± 99.00
SDSS J104433.04-012502.2	5.80	99.00	99.00 ± 99.00	46.51 ± 0.03	48.35 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J104437.06+650645.0	3.65	99.00	46.88 ± 0.03	46.60 ± 0.01	47.81 ± -1.00	4.25 ± 0.23	99.00 ± 99.00	9.57 ± 0.20	99.00 ± 99.00
SDSS J104845.05+463718.3	6.19	99.00	99.00 ± 99.00	46.55 ± 0.03	48.50 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J105123.03+354534.3	4.91	4.91	46.86 ± 0.05	46.70 ± 0.01	45.52 ± 0.18	3.35 ± 0.17	3.94 ± -1.00	9.35 ± 0.20	9.75 ± -1.00
SDSS J105756.28+455553.0	4.14	4.13	47.43 ± 0.02	47.24 ± 0.10	45.90 ± 0.09	6.93 ± 0.25	4.22 ± 0.70	10.33 ± 0.22	10.10 ± 0.23
SDSS J110657.83+081643.3	4.27	99.00	46.79 ± 0.02	46.45 ± 0.03	47.95 ± -1.00	4.85 ± 0.35	99.00 ± 99.00	9.65 ± 0.15	99.00 ± 99.00
SDSS J111348.13+540940.0	3.78	3.77	99.00 ± 99.00	46.57 ± 0.01	45.12 ± 0.30	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
BRI 1114-0822	4.49	99.00	46.58 ± 0.03	46.19 ± 0.02	48.07 ± -1.00	6.06 ± 0.12	99.00 ± 99.00	9.74 ± 0.20	99.00 ± 99.00
SDSS J113002.35+115438.3	3.40	99.00	46.91 ± 0.02	46.49 ± 0.02	47.64 ± -1.00	6.25 ± 0.08	99.00 ± 99.00	9.94 ± 0.14	99.00 ± 99.00
SDSS J113142.31+384854.6	4.13	99.00	99.00 ± 99.00	46.74 ± 0.01	47.94 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J113246.50+120901.6	5.17	99.00	99.00 ± 99.00	46.52 ± 0.03	48.10 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J113307.63+522835.5	3.74	3.74	46.71 ± 0.04	46.64 ± 0.01	45.33 ± 0.20	7.14 ± 0.37	2.50 ± -1.00	9.95 ± 0.21	9.30 ± -1.00
SDSS J113418.96+574204.6	3.52	99.00	46.98 ± 0.03	46.61 ± 0.01	47.67 ± -1.00	7.90 ± 0.84	99.00 ± 99.00	10.20 ± 0.23	99.00 ± 99.00
SDSS J114514.18+394715.9	4.06	4.03	47.02 ± 0.02	46.81 ± 0.01	45.27 ± 0.32	7.51 ± 0.34	99.00 ± 99.00	10.17 ± 0.15	99.00 ± 99.00
SDSS J114657.79+403708.6	5.01	99.00	46.54 ± 0.08	46.57 ± 0.01	48.07 ± -1.00	4.43 ± 0.71	99.00 ± 99.00	9.43 ± 0.25	99.00 ± 99.00
SDSS J114816.64+525150.2	6.42	99.00	99.00 ± 99.00	46.70 ± 0.02	48.19 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J115757.96+485655.7	4.25	99.00	46.72 ± 0.05	46.47 ± 0.01	47.93 ± -1.00	4.72 ± 0.30	99.00 ± 99.00	9.58 ± 0.21	99.00 ± 99.00
SDSS J115935.63+042420.0	3.45	99.00	46.53 ± 0.04	46.12 ± 0.03	47.62 ± -1.00	4.10 ± 0.36	99.00 ± 99.00	9.35 ± 0.16	99.00 ± 99.00

TABLE 7 — *Continued*

Name (1)	z_{ref} (2)	$z_{\text{H}\alpha}$ (3)	$\log L_{1350}$ (4)	$\log L_{5100}$ (5)	$\log L_{\text{H}\alpha}$ (6)	$\text{FWHM}_{3, \text{CIV}}$ (7)	$\text{FWHM}_{3, \text{H}\alpha}$ (8)	$\log M_{\text{BH}, \text{CIV}}$ (9)	$\log M_{\text{BH}, \text{H}\alpha}$ (10)
SDSS J120110.31+211758.5	4.58	99.00	47.00 \pm 0.04	46.63 \pm 0.01	48.03 \pm -1.00	12.09 \pm 0.48	99.00 \pm 99.00	10.60 \pm 0.22	99.00 \pm 99.00
SDSS J120207.78+323538.8	5.29	5.27	47.05 \pm 0.03	46.54 \pm 0.02	45.35 \pm 0.35	10.06 \pm 1.83	99.00 \pm 99.00	10.46 \pm 0.28	99.00 \pm 99.00
SDSS J120441.71-002149.5	5.03	99.00	46.48 \pm 0.07	46.45 \pm 0.03	47.99 \pm -1.00	5.71 \pm 1.75	99.00 \pm 99.00	9.63 \pm 0.34	99.00 \pm 99.00
SDSS J120447.15+330938.7	3.62	3.64	46.93 \pm 0.15	46.97 \pm 0.01	45.65 \pm 0.13	8.08 \pm 0.28	7.82 \pm 2.01	10.18 \pm 0.17	10.52 \pm 0.28
BR 1202-0725	4.69	4.69	47.32 \pm 0.02	46.96 \pm 0.01	45.56 \pm 0.07	9.44 \pm 0.45	4.78 \pm 0.65	10.55 \pm 0.23	10.06 \pm 0.20
SDSS J120934.54+553745.4	3.57	3.58	47.13 \pm 0.02	46.96 \pm 0.07	45.38 \pm 0.14	3.77 \pm 0.99	2.50 \pm -1.00	9.60 \pm 0.31	9.47 \pm -1.00
SDSS J122738.30+572748.9	3.99	99.00	99.00 \pm 99.00	46.79 \pm 0.01	47.93 \pm -1.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J123239.29+525250.9	4.29	4.27	47.08 \pm 0.02	46.74 \pm 0.01	45.38 \pm 0.28	7.71 \pm 0.19	99.00 \pm 99.00	10.23 \pm 0.21	99.00 \pm 99.00
SDSS J124306.55+530522.0	3.57	3.55	46.97 \pm 0.03	46.53 \pm 0.01	44.95 \pm 0.35	6.65 \pm 1.02	99.00 \pm 99.00	10.04 \pm 0.25	99.00 \pm 99.00
SDSS J125051.93+313021.9	6.13	99.00	99.00 \pm 99.00	46.46 \pm 0.01	48.45 \pm -1.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J125051.93+313021.9	6.13	6.09	99.00 \pm 99.00	46.46 \pm 0.01	44.84 \pm 0.34	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J130002.16+011823.0	4.61	99.00	46.88 \pm 0.04	46.67 \pm 0.02	48.02 \pm -1.00	5.44 \pm 0.87	99.00 \pm 99.00	9.80 \pm 0.21	99.00 \pm 99.00
SDSS J130348.94+002010.4	3.65	3.62	46.74 \pm 0.01	46.73 \pm 0.01	45.57 \pm 0.11	2.67 \pm 0.12	6.98 \pm 1.56	9.54 \pm 0.08	10.29 \pm 0.26
SDSS J130608.26+035626.3	5.99	99.00	99.00 \pm 99.00	46.36 \pm 0.11	48.47 \pm -1.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J131914.20+520200.0	3.90	3.93	47.22 \pm 0.02	46.85 \pm 0.01	45.41 \pm 0.25	4.65 \pm 0.20	99.00 \pm 99.00	9.84 \pm 0.21	99.00 \pm 99.00
SDSS J132420.83+422554.6	4.04	4.01	46.85 \pm 0.04	46.65 \pm 0.01	45.47 \pm 0.16	6.76 \pm 1.00	3.33 \pm 1.22	9.99 \pm 0.25	9.57 \pm 0.37
SDSS J132423.26+623342.1	3.63	99.00	46.48 \pm 0.06	46.35 \pm 0.01	47.68 \pm -1.00	5.69 \pm 0.13	99.00 \pm 99.00	9.63 \pm 0.20	99.00 \pm 99.00
SDSS J132603.00+295758.1	3.77	99.00	46.93 \pm 0.03	46.23 \pm 0.02	47.72 \pm -1.00	7.30 \pm 0.69	99.00 \pm 99.00	10.10 \pm 0.23	99.00 \pm 99.00
SDSS J133223.26+503431.3	3.81	99.00	47.01 \pm 0.02	46.56 \pm 0.01	47.72 \pm -1.00	6.89 \pm 1.06	99.00 \pm 99.00	10.09 \pm 0.25	99.00 \pm 99.00
SDSS J133412.56+122020.7	5.13	99.00	46.68 \pm 0.06	46.36 \pm 0.03	48.16 \pm -1.00	4.38 \pm 0.97	99.00 \pm 99.00	9.49 \pm 0.28	99.00 \pm 99.00
SDSS J133448.70+521317.9	3.61	3.60	46.71 \pm 0.05	46.51 \pm 0.01	45.17 \pm 0.25	3.68 \pm 0.25	99.00 \pm 99.00	9.35 \pm 0.20	99.00 \pm 99.00
SDSS J133529.45+410125.9	4.26	4.29	47.21 \pm 0.01	46.82 \pm 0.01	45.78 \pm 0.15	6.45 \pm 1.72	3.89 \pm 1.21	10.14 \pm 0.29	9.80 \pm 0.32
BRI 1335-0417	4.40	99.00	46.81 \pm 0.02	46.44 \pm 0.02	48.08 \pm -1.00	10.33 \pm 0.29	99.00 \pm 99.00	10.35 \pm 0.22	99.00 \pm 99.00
SDSS J134015.03+392630.7	5.03	5.05	46.68 \pm 0.06	46.48 \pm 0.02	45.15 \pm 0.30	8.94 \pm 0.53	99.00 \pm 99.00	10.15 \pm 0.22	99.00 \pm 99.00
SDSS J134040.24+281328.1	5.34	5.38	99.00 \pm 99.00	46.54 \pm 0.02	45.56 \pm 0.23	99.00 \pm 99.00	2.94 \pm 1.45	99.00 \pm 99.00	9.39 \pm 0.48
SDSS J134743.29+495621.3	4.51	4.53	47.39 \pm 0.02	46.97 \pm 0.10	45.51 \pm 0.19	7.13 \pm 1.63	6.78 \pm 2.34	10.33 \pm 0.30	10.39 \pm 0.36
BRI 1346-0322	3.99	3.99	46.64 \pm 0.03	46.82 \pm 0.01	45.26 \pm 0.16	4.77 \pm 0.08	99.00 \pm 99.00	9.55 \pm 0.19	99.00 \pm 99.00
SDSS J140132.76+411150.4	4.02	99.00	46.48 \pm 0.04	46.51 \pm 0.01	47.88 \pm -1.00	5.49 \pm 1.62	99.00 \pm 99.00	9.59 \pm 0.30	99.00 \pm 99.00
SDSS J140146.53+024434.7	4.44	99.00	47.01 \pm 0.02	46.65 \pm 0.02	48.00 \pm -1.00	6.37 \pm 2.69	99.00 \pm 99.00	10.00 \pm 0.33	99.00 \pm 99.00
SDSS J140248.08+014634.0	4.16	99.00	47.09 \pm 0.03	46.62 \pm 0.01	48.07 \pm -1.00	8.11 \pm 1.47	99.00 \pm 99.00	10.28 \pm 0.27	99.00 \pm 99.00
SDSS J141831.70+444937.5	4.31	4.33	47.04 \pm 0.02	46.54 \pm 0.01	45.32 \pm 0.29	6.10 \pm 0.39	99.00 \pm 99.00	10.00 \pm 0.16	99.00 \pm 99.00
SDSS J142144.98+351315.4	4.56	99.00	46.87 \pm 0.05	46.54 \pm 0.01	47.88 \pm -1.00	6.98 \pm 1.33	99.00 \pm 99.00	10.03 \pm 0.27	99.00 \pm 99.00
SDSS J142242.13+461310.2	3.73	99.00	46.98 \pm 0.03	46.50 \pm 0.00	47.79 \pm -1.00	7.76 \pm 0.57	99.00 \pm 99.00	10.18 \pm 0.22	99.00 \pm 99.00
SDSS J142243.02+441721.4	3.55	3.61	47.05 \pm 0.01	47.18 \pm 0.05	45.71 \pm 0.07	11.87 \pm 0.53	6.11 \pm 0.84	10.61 \pm 0.16	10.41 \pm 0.21
SDSS J143835.95+431459.2	4.61	4.67	47.35 \pm 0.01	47.14 \pm 0.05	45.39 \pm 0.20	9.84 \pm 0.42	4.92 \pm 1.89	10.60 \pm 0.14	10.19 \pm 0.39
SDSS J144144.76+472003.2	3.63	3.62	46.71 \pm 0.06	46.56 \pm 0.01	45.24 \pm 0.18	3.31 \pm 0.23	2.84 \pm 1.32	9.26 \pm 0.20	9.37 \pm 0.45
SDSS J144213.09+391856.0	3.63	99.00	47.01 \pm 0.02	46.43 \pm 0.01	47.58 \pm -1.00	7.53 \pm 0.76	99.00 \pm 99.00	10.17 \pm 0.17	99.00 \pm 99.00
SDSS J144340.70+585653.2	4.28	99.00	47.24 \pm 0.02	46.75 \pm 0.10	47.96 \pm -1.00	9.19 \pm 0.23	99.00 \pm 99.00	10.48 \pm 0.22	99.00 \pm 99.00
SDSS J144350.66+362315.1	5.27	99.00	46.78 \pm 0.03	46.58 \pm 0.01	48.08 \pm -1.00	9.46 \pm 1.49	99.00 \pm 99.00	10.25 \pm 0.21	99.00 \pm 99.00
SDSS J144542.75+490248.9	3.88	3.87	47.30 \pm 0.02	47.12 \pm 0.01	45.94 \pm 0.07	3.14 \pm 0.12	6.57 \pm 0.97	9.52 \pm 0.20	10.44 \pm 0.21
SDSS J144733.46+465723.6	3.59	99.00	46.59 \pm 0.06	46.29 \pm 0.01	47.58 \pm -1.00	4.79 \pm 0.90	99.00 \pm 99.00	9.53 \pm 0.26	99.00 \pm 99.00
SDSS J145408.95+511443.7	3.64	3.61	47.22 \pm 0.02	47.08 \pm 0.08	45.46 \pm 0.16	4.64 \pm 0.27	4.68 \pm 1.52	9.84 \pm 0.21	10.11 \pm 0.34
SDSS J150620.48+460642.4	3.50	3.52	46.84 \pm 0.02	46.38 \pm 0.01	45.21 \pm 0.22	7.21 \pm 0.21	6.98 \pm 3.02	10.03 \pm 0.15	10.11 \pm 0.42
SDSS J150654.54+522004.6	4.07	99.00	47.11 \pm 0.02	46.92 \pm 0.01	47.16 \pm -1.00	9.24 \pm 1.43	99.00 \pm 99.00	10.41 \pm 0.21	99.00 \pm 99.00
SDSS J151035.29+514841.0	5.03	5.02	99.00 \pm 99.00	46.47 \pm 0.01	45.27 \pm 0.35	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J151442.74+532412.0	3.57	99.00	46.65 \pm 0.06	46.42 \pm 0.01	47.58 \pm -1.00	5.77 \pm 0.11	99.00 \pm 99.00	9.73 \pm 0.20	99.00 \pm 99.00
SDSS J152034.53+383906.5	3.40	99.00	46.47 \pm 0.05	46.11 \pm 0.01	47.46 \pm -1.00	6.92 \pm 0.46	99.00 \pm 99.00	9.79 \pm 0.15	99.00 \pm 99.00
SDSS J152413.35+430537.4	3.93	3.95	46.83 \pm 0.02	46.58 \pm 0.01	45.23 \pm 0.29	5.07 \pm 0.89	99.00 \pm 99.00	9.71 \pm 0.21	99.00 \pm 99.00
SDSS J153650.25+500810.3	4.93	4.92	47.01 \pm 0.07	46.84 \pm 0.01	45.49 \pm 0.17	2.65 \pm 0.65	3.49 \pm 1.25	9.21 \pm 0.30	9.71 \pm 0.37
SDSS J153725.35-014650.3	3.45	99.00	46.55 \pm 0.06	46.04 \pm 0.04	47.56 \pm -1.00	5.62 \pm 0.15	99.00 \pm 99.00	9.65 \pm 0.20	99.00 \pm 99.00
SDSS J153825.74+420933.3	3.83	99.00	46.82 \pm 0.02	46.52 \pm 0.01	47.79 \pm -1.00	7.78 \pm 0.60	99.00 \pm 99.00	10.10 \pm 0.16	99.00 \pm 99.00
SDSS J154340.38+341744.4	4.41	99.00	47.15 \pm 0.02	47.01 \pm 0.01	48.02 \pm -1.00	5.16 \pm 0.15	99.00 \pm 99.00	9.90 \pm 0.15	99.00 \pm 99.00
SDSS J154914.45+364822.0	3.58	3.60	46.59 \pm 0.03	46.26 \pm 0.01	44.88 \pm 0.35	6.96 \pm 0.51	99.00 \pm 99.00	9.87 \pm 0.16	99.00 \pm 99.00
SDSS J160254.18+422822.9	6.07	99.00	99.00 \pm 99.00	46.65 \pm 0.01	48.47 \pm -1.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J161140.13+273029.6	3.33	99.00	46.70 \pm 0.04	46.30 \pm 0.02	47.58 \pm -1.00	4.52 \pm 0.38	99.00 \pm 99.00	9.53 \pm 0.16	99.00 \pm 99.00
SDSS J161425.13+464028.9	5.31	99.00	99.00 \pm 99.00	46.59 \pm 0.01	48.07 \pm -1.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00	99.00 \pm 99.00
SDSS J162100.70+515544.8	5.59	5.59	99.00 \pm 99.00	46.72 \pm 0.01	45.20 \pm 0.18	99.00 \pm 99.00	3.09 \pm 1.17	99.00 \pm 99.00	9.53 \pm 0.38
SDSS J162331.81+311200.5	6.22	6.23	99.00 \pm 99.00	46.39 \pm 0.04	44.83 \pm 0.21	99.00 \pm 99.00	4.37 \pm -1.00	99.00 \pm 99.00	9.68 \pm -1.00
SDSS J162520.31+225832.9	3.77	3.78	47.11 \pm 0.01	46.66 \pm 0.06	45.44 \pm 0.17	5.02 \pm 0.18	5.37 \pm 1.80	9.85 \pm 0.15	10.01 \pm 0.34
SDSS J162623.38+484136.4									

TABLE 7 — *Continued*

Name (1)	z_{ref} (2)	$z_{\text{H}\alpha}$ (3)	$\log L_{1350}$ (4)	$\log L_{5100}$ (5)	$\log L_{\text{H}\alpha}$ (6)	$\text{FWHM}_{3, \text{C IV}}$ (7)	$\text{FWHM}_{3, \text{H}\alpha}$ (8)	$\log M_{\text{BH}, \text{C IV}}$ (9)	$\log M_{\text{BH}, \text{H}\alpha}$ (10)
BR 2248-1242	4.16	4.14	46.66 ± 0.07	46.81 ± 0.01	45.82 ± 0.10	1.58 ± 0.07	4.74 ± 0.96	8.55 ± 0.19	9.98 ± 0.24
BR J2328-4513	4.36	99.00	99.00 ± 99.00	46.41 ± 0.01	47.92 ± -1.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00	99.00 ± 99.00
SDSS J233330.17+152538.7	3.68	3.68	46.61 ± 0.06	46.31 ± 0.02	45.19 ± 0.26	6.07 ± 0.88	99.00 ± 99.00	9.75 ± 0.24	99.00 ± 99.00
BR J2349-3712	4.21	99.00	46.86 ± 0.02	46.53 ± 0.01	47.95 ± -1.00	5.73 ± 0.37	99.00 ± 99.00	9.84 ± 0.21	99.00 ± 99.00

NOTE. — Catalog of the properties derived for the *AKARI* quasars, sorted by right ascension. Column 1: Target name; Column 2: Redshift from references; Column 3: Redshift measured from $\text{H}\alpha$; Column 4: 1350 Å luminosity and its uncertainty; Column 5: 5100 Å luminosity and its uncertainty; Column 6: $\text{H}\alpha$ luminosity and its uncertainty; Column 7: FWHM of the C iv line and its uncertainty; Column 8: FWHM of the $\text{H}\alpha$ line and its uncertainty; Column 9: M_{BH} from the C iv line and its uncertainty; Column 10: M_{BH} from the $\text{H}\alpha$ line and its uncertainty. The units for L , FWHM, and M_{BH} are ergs s^{-1} , 1000 km s^{-1} , and M_{\odot} . Columns 9 and 10 are from Equations (10) and (7), respectively. Empty parameters are entered as 99 and upper limits are given with errors of -1.